

Euclid Quick Data Release (Q1)

XXXVIII. *Euclid* spectroscopy of quasars. 1. Identification and redshift determination of 3500 bright quasars

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ABSTRACT

The slitless spectroscopy mode of the Near-Infrared Spectrometer and Photometer (NISP) on board the *Euclid* telescope has enabled efficient spectroscopy of objects within a large field of view. Nevertheless, the relatively low spectral resolution, overlapping spectra, and contamination pose challenges to source classification and redshift determination using the NISP spectra alone. In this work, we present a large and homogeneous sample of bright quasars identified from the *Euclid* Quick Data Release (Q1), constructed by combining high-purity candidate selections from *Gaia* and WISE with the new spectroscopic capabilities of *Euclid*. Through visual inspection of the *Euclid* spectra of these quasar candidates, we identify approximately 3500 quasars and determine reliable redshifts in the range of $0 < z \leq 4.8$. Of these, 2686 are new spectroscopic identifications relative to existing public compilations. We generated the first *Euclid* composite spectrum of quasars covering rest-frame near-ultraviolet (NUV) to near-infrared (NIR) wavelengths without telluric lines, which will be pivotal to NIR quasar spectral analysis. We obtained an empirical spectroscopic depth of $J_E \lesssim 21.5$ and $H_E \lesssim 21.3$ at the sensitivity of the Wide Field Survey, beyond which the number of securely identified quasars declines sharply. Accordingly, the sample presented in this paper comprises spectroscopically confirmed quasars brighter than these limits. We analysed morphological parameters from the Visible Camera (VIS) using Sérsic and model-independent (CAS) metrics, and a deep-learning point spread function fraction (f_{PSF}) to track nuclear dominance. The VIS morphologies show a clear redshift dependence: at low redshift ($z < 0.5$), obvious host structures are common and a single Sérsic model fits about half of the sources; at intermediate redshift ($0.5 < z < 2$), the nuclear component dominates, with 90% of the Sérsic fits saturating at the upper index limit. In this intermediate redshift regime, f_{PSF} is available, and we use it as a more reliable compactness measure than the single-Sérsic and CAS parameters to quantify nuclear versus host emission. We also explore the novel *Euclid* NIR colour space and discuss the role of these quasars in refining active galactic nucleus selection techniques for future *Euclid* data releases.

Key words. techniques: spectroscopic – catalogs – galaxies: active – galaxies: distances and redshifts – quasars: general – infrared: galaxies

1. Introduction

Powered by accreting supermassive black holes, quasars are among the most luminous and distant objects in the Universe. Quasars act as beacons that allow us to probe recent to early cosmic epochs and trace the large-scale structure of the cosmos (e.g. Blanton et al. 2017; Neveux et al. 2020; Fan et al. 2023). Traditionally, quasar surveys have primarily relied on multi-colour selection techniques to isolate candidates, with subsequent slit or multi-fibre spectroscopy to confirm their nature and determine precise redshifts (e.g. Richards et al. 2002; Croom et al. 2004; Myers et al. 2015; Chaussidon et al. 2023). Ground-based quasar surveys have adopted slitless spectroscopy to find quasars in a cost-effective way. However, such slitless campaigns suffer from low spectral resolution, overlapping spectra, and contamination, often necessitating follow-up slit spectroscopy for confirmation (e.g. Schmidt et al. 1986; Osmer & Hewett 1991; Schneider et al. 1999).

The advent of space-based observations and improvements in data reduction have substantially improved the quality of slitless spectroscopic data. For example, the *Hubble* Space Telescope's (HST) Advanced Camera for Surveys (ACS) and Wide Field Camera 3 (WFC3) have produced high-quality grism data (e.g. Momcheva et al. 2016; Estrada-Carpenter et al. 2019)

that are processed with dedicated pipelines (e.g. Kümmel et al. 2009; Brammer 2019). More recently, with the low background, high sensitivity, and high spatial resolution of the *James Webb* Space Telescope (JWST), the Near Infrared Camera (NIRCam; Rieke et al. 2005, 2023) and the Near Infrared Imager and Slitless Spectrograph (NIRISS; Willott et al. 2022; Doyon et al. 2023) on board JWST have enabled unprecedented studies of distant galaxies (e.g. Roberts-Borsani et al. 2022; Sun et al. 2023; Oesch et al. 2023; Meyer et al. 2024).

The European Space Agency's (ESA) *Euclid* mission (Euclid Collaboration: Mellier et al. 2025) is designed to probe the dark matter and dark energy of the Universe by studying weak lensing and galaxy clustering over approximately one third of the sky in both the optical and near-infrared using the Visible Camera (VIS; Euclid Collaboration: Cropper et al. 2025) and the Near-Infrared Spectrometer and Photometer (NISP; Euclid Collaboration: Jahnke et al. 2025). A key feature of *Euclid* is the slitless spectroscopic mode of NISP, which can simultaneously capture spectra of sources in a large field of view of 0.57 deg^2 . Three red grisms covering the same R_{GE} band (1206–1892 nm) are adopted in the Euclid Wide Survey to provide spectra with different dispersion directions of 0° , 180° , and 270° with respect to the detector columns. Dispersed slitless images of the three grisms are combined to disentangle

overlapping spectra from multiple sources, and generate the final clean spectra. The red grisms have a resolving power of $\mathcal{R}_{\text{RGE}} > 480$ for a source with a $0''.5$ diameter (Euclid Collaboration: Jahnke et al. 2025; Euclid Collaboration: Mellier et al. 2025). This spectroscopic capability enables a census of bright quasars and enhances the efficiency of quasar discovery by leveraging the high sensitivity and spatial resolution from space. For example, Bañados et al. (2025) have recently discovered a $z = 5.404$ quasar, EUCL J181530.01+652054.0, using NISP spectroscopy.

Euclid Collaboration: Lusso et al. (2024) provide a detailed prediction of the NISP spectroscopic mode for active galactic nuclei (AGNs) using mock spectra. They demonstrate that redshift measurements are robust when the $\text{H}\alpha$ emission line is visible within the spectral coverage of RGE ($0.89 < z < 1.83$) at a line flux greater than $2 \times 10^{-16} \text{ erg s}^{-1} \text{ cm}^{-2}$. Outside this redshift range, however, redshift measurements are inefficient due to a low signal-to-noise ratio (S/N) or a lack of prominent emission lines or both.

Early investigations of AGNs in the first *Euclid* Quick Data Release (Q1; Euclid Collaboration: Aussel et al. 2026) already demonstrate the potential of *Euclid* for AGN science. For example, Euclid Collaboration: Matamoro Zatarain et al. (2026) present an AGN candidate catalogue with a total of 229 779 objects selected with multi-wavelength data, while Euclid Collaboration: Roster et al. (2026) identify *Euclid* counterparts to X-ray sources in the Deep Fields, most of which are AGNs, using catalogues from *eROSITA* (Merloni et al. 2024), *XMM-Newton* (Webb et al. 2020), and *Chandra* (Evans et al. 2024).

In this work, we present a homogenous bright quasar sample identified with spectroscopic data of Q1, complementary to the efforts of Euclid Collaboration: Matamoro Zatarain et al. (2026) and Euclid Collaboration: Roster et al. (2026). We select quasar candidates from all-sky optical and mid-infrared databases, and identify the objects based on prominent emission lines in the *Euclid* slitless spectra. Our analysis demonstrates that slitless spectroscopy, as implemented in the *Euclid* mission, not only overcomes some of the limitations of colour-based quasar selection techniques, but also provides a valuable dataset for studying the large-scale structure of the Universe through quasar clustering and cross-correlation with galaxy and weak lensing maps (e.g. Myers et al. 2003; Pullen et al. 2015; Petter et al. 2023; Alonso et al. 2023). Spectral properties of these quasars obtained with multi-component fitting, including spectral indices, line widths, and black hole masses, will be reported in Euclid Collaboration: Calhau et al. (in preparation).

The paper is organised as follows. Section 2 describes the sample selections of quasar candidates and the crossmatch to the *Euclid* spectral catalogue. Section 3 describes the identification and redshift determination procedure. We present our results in Sect. 4 and discuss their implications in Sect. 5. Finally, Sect. 6 concludes with a summary of our findings and prospects for future studies. All magnitudes are in the AB system (Oke & Gunn 1983) unless stated otherwise.

2. Data

In this work, we use external quasar candidates from *Gaia* and AllWISE as the input sample for the identification with *Euclid* spectroscopy. Below, we briefly introduce the Q1 data we use, and describe the contents of the input quasar candidate sample and the matched Q1 spectroscopic sample.

2.1. Q1 photometry and spectroscopy

The Q1 (Euclid Collaboration: Aussel et al. 2026; Euclid Quick Release Q1 2025) dataset contains images and photometric catalogues from both VIS (I_E band) and NISP (Y_E , J_E , and H_E bands), and one-dimensional (1D) spectra of the NISP spectroscopic mode. In this work, we use the main photometric catalogue (catalogue.mer_catalogue in the *Euclid* Science Archive¹) generated by the MERge Processing Function (MER; Euclid Collaboration: Romelli et al. 2026), which include aperture fluxes, template-fit and Sérsic-fit fluxes, and quality flags in each band, as well as morphological information for all sources detected in the Euclid Deep Fields. We also use the main morphology catalogue (catalogue.mer_morphology; Euclid Collaboration: Quilley et al. 2026) to assess the morphological properties of the final sample.

Combined 1D slitless spectroscopic data of sources brighter than $H_E = 22.5$ have been generated by the dedicated SIR spectroscopic processing function (Euclid Collaboration: Copin et al. 2026). The slitless spectra used in this work are retrieved from ESA Datalabs² (Navarro et al. 2024).

Because the observed slitless spectrum of an object is the convolution of its intrinsic spectrum with its wavelength-dependent spatial light profile along the dispersion direction, sources with larger extents will have lower spectral resolution. In addition, the blue and red ends of the observed spectrum tend to show upturns due to the convolution nature of the observed spectrum, and the lower S/N at both ends (see e.g. Pirzkal et al. 2004). Each of the original 1D spectra contains 531 data points, covering 11 900–19 002 Å with a wavelength interval of 13.4 Å. To keep only the useful data from the slitless spectra, we trim the blue and red ends, retaining 12 047–18 734 Å, which yields 500 data points per spectrum.

2.2. Reliable quasar candidates from *Gaia* and AllWISE

Gaia DR3 announced a sample of 6.6 million quasar candidates (the qso_candidates table, hereafter the GDR3 QSO candidate catalogue; Gaia Collaboration 2023a,b), which has high completeness thanks to the combination of several different modules, including the Discrete Source Classifier (DSC), the Quasar Classifier (QSOC), the variability classification module, the surface brightness profile module, and the *Gaia* DR3 Celestial Reference Frame source table. The DSC uses the *Gaia* Blue (BP) and Red (RP) Photometer (De Angeli et al. 2023) spectrum together with the mean *G*-band magnitude, the variability in this band, the parallax, and the proper motion to classify each *Gaia* source probabilistically into five classes: quasars, galaxies, stars, white dwarfs, and physical binary stars. The QSOC determines the redshifts using BP and RP spectra of the sources classified as quasars by the DSC (Delchambre et al. 2023). The variability classification module identifies 25 classes of variable sources (including AGN candidates) from the variability of the *Gaia* light curves using supervised machine learning (Rimoldini et al. 2023; Carnerero et al. 2023). Both the surface-brightness profile module and the *Gaia* DR3 Celestial Reference Frame source table are based on external catalogues of quasars and quasar candidates (see Ducourant et al. 2023; Gaia Collaboration 2022, for a complete list). Despite its high

¹ <https://eas.esac.esa.int/sas/>

² <https://datalabs.esa.int/>

completeness, the GDR3 QSO candidate catalogue has an estimated low purity of quasars (52%) and a large scatter of redshift estimates (Gaia Collaboration 2023b).

Instead of using the original GDR3 QSO candidate catalogue, we take its purified subsets to find *Euclid* counterparts of the sources. These purified catalogues include the following.

1. Quaia (Storey-Fisher et al. 2024) with nearly 1.3 million sources at $G < 20.5$. This sample is selected using a set of cuts involving proper motion, *Gaia* and UnWISE (Schlafly et al. 2019) colours and magnitudes, which are designed to remove stellar contaminants of the Milky Way and the Large and Small Magellanic Clouds.
2. CatNorth (Fu et al. 2024) with more than 1.5 million sources down to the *Gaia* limiting magnitude ($G < 21.0$) in the 3π sky of the Pan-STARRS1 (PS1; Chambers et al. 2016) footprint ($\delta > -30^\circ$). This catalogue is primarily built with a machine learning classification model trained on colour and morphological features from *Gaia*, PS1, and CatWISE2020 (Marocco et al. 2021). An additional probabilistic cut on proper motion (probability density of zero proper motion; see Fu et al. 2021, 2024, for the definition) is applied to further purify the candidates. The CatNorth catalogue has a purity of approximately 90%.
3. CatSouth (Fu et al. 2025) with 0.9 million sources with $G < 21.0$ covered by the fourth data release (DR4) of the SkyMapper Southern Survey (SMSS; $\delta \lesssim 16^\circ$; Onken et al. 2024). This catalogue is built with the same method as CatNorth, while based on data from *Gaia*, SMSS DR4, and VISTA (Visible and Infrared Survey Telescope for Astronomy) surveys (Emerson et al. 2006; Minniti et al. 2010; Cioni et al. 2011; McMahon et al. 2013; Edge et al. 2013), as well as CatWISE2020.

All three purified catalogues propagate the *Gaia* DR3 QSO template-matching redshift and combine it with multi-band photometry to derive improved photometric redshifts. These photometric redshifts reduce the fractions of catastrophic outliers by more than 15% compared to the original QSO redshift, at the cost of a modest decrease in precision (Storey-Fisher et al. 2024; Fu et al. 2024). The compilation of the three catalogues above contains more than 1.9 million unique (with unique *Gaia* source_id) quasar candidates in the entire sky. This purified GDR3 QSO candidate catalogue will be referred to as GDR3-QSOs hereafter for simplicity. Together, 4467 sources are matched to the Q1 main photometric catalogue (see Sect. 2.1 for details) using *Gaia* source_id, which is listed in both GDR3-QSOs and catalogue_mer_catalogue of Q1.

In addition to GDR3-QSOs, which are mainly optically bright quasars, we also include AGN candidates selected with data from the Wide-field Infrared Survey Explorer (WISE; Wright et al. 2010), a NASA mission that has surveyed the entire sky in the 3.4-, 4.6-, 12-, and 22- μ m mid-infrared bands (W1, W2, W3, and W4). The AllWISE source catalogue was built by combining data from the WISE cryogenic and NEOWISE (Mainzer et al. 2011) post-cryogenic survey phases, providing positions, proper motions, four-band fluxes, and flux variability statistics for over 747 million objects. Using W1 – W2 colour and W2 magnitude from AllWISE, Assef et al. (2018) constructed two large catalogues of AGN candidates across 75% of the sky: the R90 catalogue with 90% reliability, and the C75 catalogue with 75% completeness. In total, 8202 sources from the R90 AGN candidate catalogue are matched to the Q1 main photometric catalogue using a radius of $1''.5$.

Combining GDR3-QSOs and the R90 AGN candidates yields 10 201 unique Q1 sources. Among them, 5083 are

Table 1. Summary of numbers of input sources and identified quasars.

Sample	<i>Gaia</i>	Non- <i>Gaia</i>	All
Photometric (MER) sources	5083	5118	10 201
Spectroscopic (SIR) sources	4965	4249	9214
Identified quasars	2753	715	3468
Success rate	55.4%	16.8%	37.6%

Notes. The success rate is the number of identified quasars divided by that of SIR sources.

detected by *Gaia* (including 616 sources from R90), and 5118 are not in *Gaia*. We refer to these two subsets as *Gaia* and non-*Gaia* subsets, the latter only selected using the R90 AGN candidate catalogue. Matching the input quasar candidates with the Q1 spectra source table (sedm.spectra_source in the *Euclid* Science Archive) gives 9214 sources. The numbers of input sources of different subsets are summarised in Table 1.

3. Source classification and spectral redshift determination

3.1. Template-matching redshifts for the candidates

For each observed quasar spectrum, we first estimated the redshift by comparing it to a rest-frame template via a Pearson correlation function (PCF, also known as the normalised cross-correlation) method (Sartoretti et al. 2018). We began with the template from Glikman et al. (2006), based on 27 bright quasars observed with the NASA Infrared Telescope Facility (IRTF). After constructing a new composite spectrum in Sect. 4.3, we re-ran the PCF with the new composite as a template to refine the redshift estimates; for the final measurements we adopted a piecewise template that combines Vanden Berk et al. (2001), the mean composite from this work, and Glikman et al. (2006), see Appendix C.2. The PCF algorithm proceeds as follows:

1. The template wavelengths are shifted into the observed frame using the relation

$$\lambda_{\text{obs}} = \lambda_{\text{rest}} (1 + z), \quad (1)$$

where z is the trial redshift.

2. The shifted template is then interpolated onto the observed wavelength grid. Both the observed flux, $F_{\text{obs}}(\lambda)$, and the interpolated template flux, $F_{\text{temp}}(\lambda; z)$, are median-normalised as

$$\tilde{F}_{\text{obs}}(\lambda) = \frac{F_{\text{obs}}(\lambda)}{\text{median}[F_{\text{obs}}(\lambda)]}, \quad (2)$$

$$\tilde{F}_{\text{temp}}(\lambda; z) = \frac{F_{\text{temp}}(\lambda; z)}{\text{median}[F_{\text{temp}}(\lambda; z)]}, \quad (3)$$

to minimise the effect of continuum differences.

3. The PCF is computed over the set, $\mathcal{I}$, of overlapping wavelength bins using the median-normalised spectra above:

$$r(z) = \frac{\sum_{i \in \mathcal{I}} [\tilde{F}_{\text{obs},i} - \langle \tilde{F}_{\text{obs}} \rangle] [\tilde{F}_{\text{temp},i}(z) - \langle \tilde{F}_{\text{temp}}(z) \rangle]}{\sqrt{\sum_{i \in \mathcal{I}} [\tilde{F}_{\text{obs},i} - \langle \tilde{F}_{\text{obs}} \rangle]^2 \sum_{i \in \mathcal{I}} [\tilde{F}_{\text{temp},i}(z) - \langle \tilde{F}_{\text{temp}}(z) \rangle]^2}}, \quad (4)$$

where $\langle \tilde{F} \rangle$ denotes the mean of the median-normalised spectrum.

4. To account for the logarithmic nature of wavelength shifts, the redshift grid is sampled with a step size that scales with $1 + z$:

$$\Delta z = \delta (1 + z), \quad (5)$$

where δ is a base step (e.g. $\delta = 0.001$ in this work).

5. The best-fit template redshift (hereafter z_{temp}) is then adopted as the value of z that maximises $r(z)$:

$$z_{\text{temp}} = \arg \max_z r(z). \quad (6)$$

A primary advantage of using the PCF is that it constrains the correlation to be in the range between -1 and $+1$, thereby removing dependence on the absolute flux scale or any additive offsets between the observed spectrum and the template. This normalisation ensures that the metric is driven solely by the relative shapes and positions of spectral features. By employing the PCF method, our analysis emphasises the similarity in spectral features rather than overall flux levels.

3.2. Visual inspection

An interactive visual inspection tool PGSpecPlot, part of the Python package spechbox (Fu 2026), was used to check the spectra of the quasar candidates. During visual inspection, each spectrum is displayed sequentially with an overplotted quasar template adjusted to the estimated template redshift (z_{temp}). Using interactive controls, including a slider with non-linear $(1 + z)$ scaling and a corresponding spin box (a text box with an up-down control), the user can verify the template redshift and adjust it as necessary. Keyboard shortcuts facilitate the rapid classification of each spectrum (e.g. flagging non-quasar objects or uncertain cases). A history of inspected spectra is maintained in a comma-separated value file so that previously processed spectra can be automatically loaded.

In addition to the initial z_{temp} , a *Gaia* redshift is displayed in the window when available. The *Gaia* redshift (hereafter z_{Gaia}) is primarily based on the *Gaia* DR3 low-resolution spectral template-matching redshift from QSOC (redshift_qsoc; Delchambre et al. 2023; Gaia Collaboration 2023b), supplemented by the photometric redshift from CatNorth and CatSouth when redshift_qsoc is not available. Because z_{Gaia} and z_{temp} are independent spectroscopic redshift estimates obtained from different wavelength ranges (z_{Gaia} from the BP or RP bands within 330–1050 nm, and z_{temp} from RG_E in the NIR), a z_{temp} that is close to z_{Gaia} (with $|z_{\text{temp}} - z_{\text{Gaia}}|/[1 + z_{\text{temp}}] < 0.15$) is taken as a secure visual redshift (z_{vi}) in most cases, even when only one emission line is present in the wavelength range of RG_E. Nevertheless, when z_{Gaia} is unavailable and only one emission line is seen, the redshift can be highly uncertain. Such single-line spectra are labelled as ‘uncertain-redshift’ objects unless the emission line and continuum fitted the template with high confidence during visual inspection.

As expected from the predictions (Euclid Collaboration: Lusso et al. 2024), quasars with $0.89 < z < 1.83$ are most easily identified with the H α emission line, which is the strongest and broadest line among all emission lines detected by RG_E. Quasars at lower redshift ($z < 0.89$) are mainly identified with the combination of several rest-frame NIR emission lines, i.e. [S III] $\lambda 9071$, Pa δ , He I + Pa γ (blended), and Pa β . At $1.83 < z < 2.85$, quasars are identified with H β , [O III], and H γ . At $2.85 < z < 3.3$, the quasar spectra lack strong features except for H γ and the pseudo-continuum

(‘small blue bump’), and they are mainly identified with the agreement between z_{temp} and z_{Gaia} . At $z > 3.3$, Mg II enters the wavelength range, and the redshift determination is secured with the combination of z_{temp} (from the Mg II line) and z_{Gaia} . Figure 1 shows example *Euclid* spectra of identified quasars at different rest-frame wavelength ranges, which correspond to the different emission line features described above.

Approximately 2300 sources are labelled as ‘uncertain-redshift’ objects, making up 25% of the entire sample (9214 spectra). Given the complexity of the data reduction and limited depth of the slitless spectroscopy, there are also approximately 3400 unidentifiable spectra with: (i) corrupted data with too many invalid or anomalously high flux values; or (ii) featureless spectra due to either weak lines or low S/N. These spectra are currently classified as unknown sources. However, given the high reliability of our input catalogues, it is likely that many of these objects are indeed quasars that could be confirmed with deeper observations or improved data processing in the future.

3.3. Spurious redshift rejection, consistency check with DESI, and the final redshift

To evaluate the performance of our visual redshift z_{vi} , we first examined the source concentration parameter $\mu_{\text{max}} - \text{mag}$ (mumax_minus_mag in catalogue_mer_catalogue). Here, μ_{max} is the source peak surface brightness above the background, and mag is the magnitude used to compute point-like probability, both given by SourceXtractor++³ (Bertin et al. 2020; Kümmel et al. 2022) during the MER data reduction. The estimator $\mu_{\text{max}} - \text{mag}$ is related to the concentration of light at the peak versus the total magnitude; at a given magnitude, sources with smaller $\mu_{\text{max}} - \text{mag}$ are more point-like (Euclid Collaboration: Romelli et al. 2026). This parameter has also been used in Jauzac et al. (2012), Sharon et al. (2022), Estrada et al. (2023), Euclid Collaboration: Matamoros Zatarain et al. (2026), and Euclid Collaboration: Roster et al. (2026) as input for point/extended source classification.

Figure 2 shows the distribution of visually identified quasars in the $\mu_{\text{max}} - \text{mag} - z_{\text{vi}}$ plane. In general, $\mu_{\text{max}} - \text{mag}$ decreases as z_{vi} becomes higher. Most sources exhibit compact morphologies (low $\mu_{\text{max}} - \text{mag}$ values) at $z_{\text{vi}} \gtrsim 0.8$, consistent with unresolved point-like sources. However, a distinct subset of sources at $z_{\text{vi}} > 2$ shows significantly higher $\mu_{\text{max}} - \text{mag}$ values, indicative of extended or low-concentration profiles that are atypical for high-redshift quasars. These sources are flagged as spurious redshifts and excluded from the final sample using a conservative empirical cut: $\mu_{\text{max}} - \text{mag} > -2$ at $z_{\text{vi}} > 2$, as indicated by the dashed red lines. This selection removes sources whose morphology and redshift are inconsistent with expectations for high-redshift quasars, improving the redshift accuracy of the final catalogue.

As an external consistency check on our visually confirmed redshifts, we crossmatched our quasar sample with Data Release 1 of the Dark Energy Spectroscopic Instrument (DESI DR1; DESI Collaboration 2025) using a radius of $1''.0$ and obtained 454 objects in common. We flag discrepant cases using the criterion $|z_{\text{vi}} - z_{\text{DESI}}|/(1 + z_{\text{DESI}}) > 0.15$, yielding 16 outliers. We then visually reinspect both the DESI and *Euclid* spectra for all 16 sources. In 11 cases, the DESI spectra support the same redshift solution as our *Euclid*-based identification, while the DESI catalogue redshift is inconsistent with the spectral fea-

³ <https://github.com/astrorama/SourceXtractorPlusPlus>

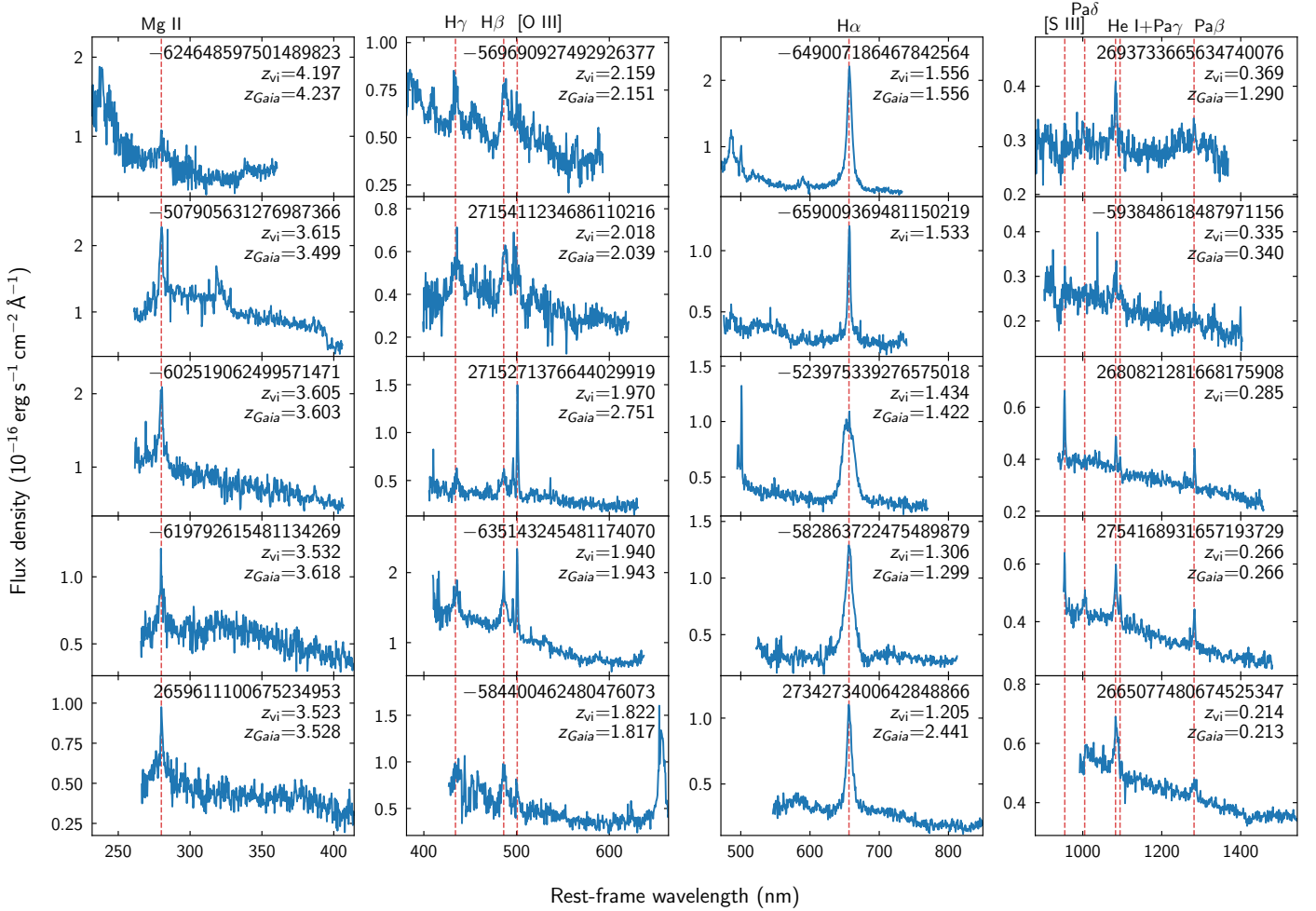


Fig. 1. Example *Euclid* spectra of the visually identified quasars in the rest frame. The four columns from left to right show prominent emission lines used for the visual inspection in descending order of redshift: Mg II in the first column; H γ , H β , [O III], and H α in the second column; H α in the third column; and He I + Pa γ , Pa β , Pa δ , and [S III] in the last column.

tures. In the remaining five cases, we revise our visual redshifts; the updated z_{vi} values are now consistent with z_{DESI} and are adopted in the released catalogue. The spectra of the 16 initially discrepant sources are shown in Appendix A.

The visually identified quasar sample contains 3468 sources (38% of the entire sample) covering the redshift range of $0 < z \lesssim 4.8$. In addition to the crossmatch to DESI DR1 above, we also crossmatch the full sample with the Million Quasar Catalogue (Milliquas v8; [Flesch 2023](#)) using a radius of $1''.0$ and find 710 sources in common. The union of the DESI DR1 matches (454 objects) and the Milliquas matches (710 objects) contain 782 sources, implying that 2686 quasars in our sample are new spectroscopic identifications.

4. The bright quasar sample identified with Q1 spectroscopy

4.1. Photometric properties

The 3468 sources of the visually identified quasar sample are compiled into a catalogue detailed in Table B.1. This catalogue includes source IDs and coordinates, redshifts, spectroscopic quality indicators, point spread function (PSF) fraction measurements from VIS imaging (*Euclid* Collaboration: [Margalef-Bentabol et al. 2026](#)) for a redshift-limited subsample ($0.5 < z < 2$), and magnitudes across

the *Euclid* bands. Among them, 2753 are *Gaia* sources and 715 are non-*Gaia* sources.

The redshift and magnitude distributions of the full sample, and *Gaia* or non-*Gaia* subsets, are shown in Fig. 3. The redshift density distributions peak in the range $0.89 \lesssim z \lesssim 1.83$ when the H α emission line is in the observed wavelength range, and drop sharply at $z \approx 0.89$ and $z \approx 1.83$ when H α moves out of the wavelength range. As shown in the histograms of *Euclid* magnitudes, these visually identified quasars represent a bright sample with a median I_E of 20.5, and median magnitudes of 19.9 in Y_E , 19.7 in J_E , and 19.5 in H_E . The non-*Gaia* subset is overall 2 magnitudes fainter than the *Gaia* subset in I_E , 0.8 magnitudes fainter in J_E and 0.6 magnitudes fainter in H_E . The faintest identified quasars in this work have $I_E \approx 27$, $Y_E \approx 23$, and $J_E \approx H_E \approx 22.5$.

To assess the effective depth of reliable spectral identification in our sample, we examined the relationship between the median S/N of the NISP spectra within $[12\,047, 18\,734] \text{ Å}$ and the J_E and H_E magnitudes. As shown in Fig. 4, the number of visually confirmed quasars declines steeply below $S/N = 2$, indicating an empirical limit where spectral identifications become increasingly difficult. By fitting a linear relation to $\log_{10}(S/N)$ as a function of magnitude, we find that this empirical transition occurs at approximately $J_E = 21.5$ and $H_E = 21.3$. These values define the practical magnitude limits beyond which reliable redshift determination becomes rare in Q1 slitless spectroscopy for quasars.

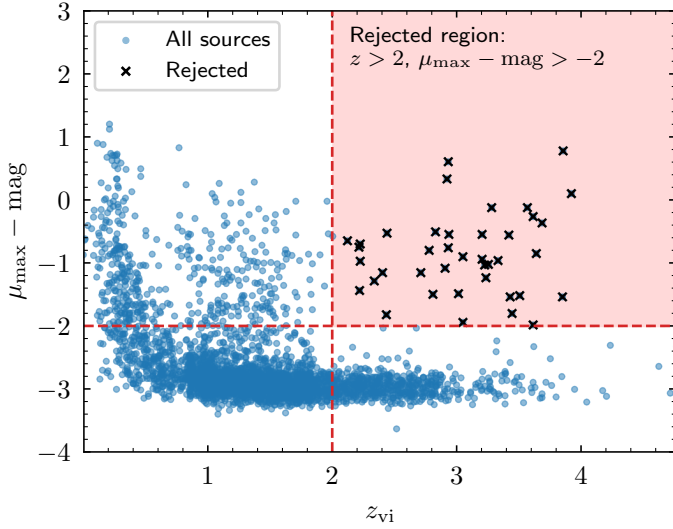


Fig. 2. Concentration parameter $\mu_{\max} - \text{mag}$ as a function of the visual redshift z_{vi} . The rejected region is shaded in pink, and is defined with $\mu_{\max} - \text{mag} > -2$ at $z_{\text{vi}} > 2$, as indicated by the dashed red lines. Sources inside the rejected region are considered to have spurious redshifts and are marked with black crosses.

We also compare the colour distributions of the *Gaia* and non-*Gaia* quasar subsets on the AllWISE and *Euclid* colour-colour diagrams. To illustrate colour spaces occupied by stars, we show in the diagrams *Euclid* point-like sources that are selected from catalogue.mer_catalogue using

```
mumax_minus_mag < -2.6 AND spurious_flag = 0
AND flux_vis_psf > 0 AND flux_y_tmplfit > 0
AND flux_h_tmplfit > 0 AND flux_j_tmplfit > 0.
```

These point-like criteria are similar to the definition of *Euclid* point-like sources in Euclid Collaboration: Matamoro Zatarain et al. (2026).

The *Gaia* and non-*Gaia* quasar subsets share nearly identical colour spaces on the $W1 - W2$ versus $W2 - W3$ diagram (Fig. 5). These have been described as AGN regions by many previous studies (e.g. Stern et al. 2012; Wu et al. 2012; Mateos et al. 2012; Assef et al. 2018). On *Euclid* colour planes (Fig. 6), however, the *Gaia* and non-*Gaia* quasar samples become more separable. The non-*Gaia* subset occupies colour spaces redder than the *Gaia* subset, while partly overlapping. The optical-faint and infrared-bright selection of the non-*Gaia* subset makes it a representative sample of red quasars. As shown in Fig. 6, the *Euclid* colour cuts for red quasars ($Y_E - H_E > 0.7$, $J_E - H_E > 0.3$, and $I_E - H_E > 1.8$) from Euclid Collaboration: Tarsitano et al. (2026, hereafter T25) select the redder half of non-*Gaia* quasars, which only have a small overlap with the *Gaia* quasars.

4.2. A stacked emission-line map and the redshift challenge

To ensure robust composite building and follow-up spectral analysis, we selected a golden sample of 2868 visually identified quasars with the following constraints: (i) median S/N of the spectrum higher than 3 within $[12047, 18734] \text{ \AA}$; and (ii) containing ≤ 15 invalid pixels (NaN, or zero flux values).

To better understand the systematic effects of the redshift determination using RG_E spectra, we constructed a stacked emission-line map of the golden sample on the redshift-wavelength plane (Fig. 7) after normalising each spectrum with

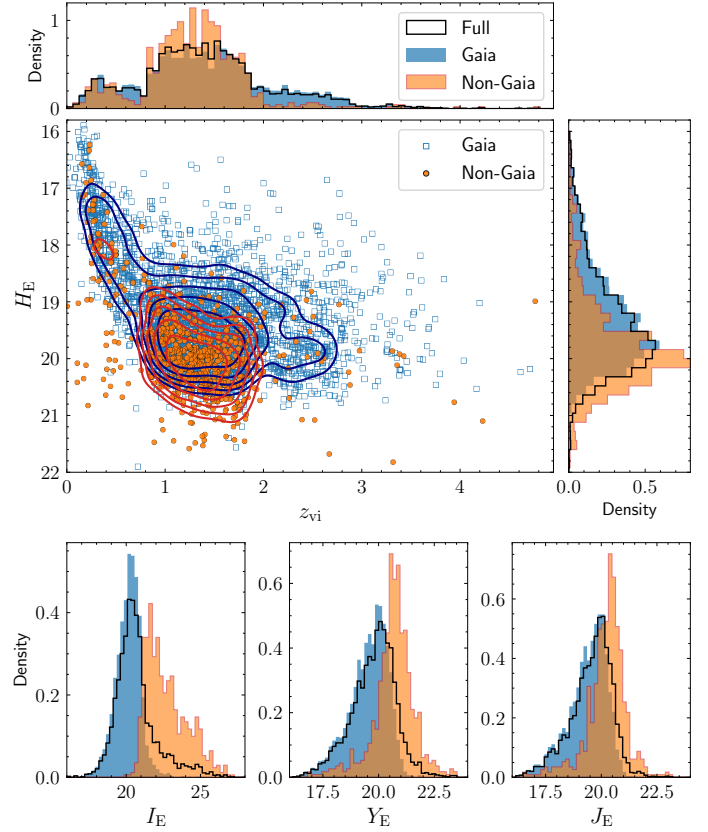


Fig. 3. Redshift-magnitude distributions for the visually identified quasar sample using *Euclid* Q1 data. *Top*: Two-dimensional distribution of H_E versus z_{vi} . *Gaia*-detected sources are shown as blue open squares, and sources not in *Gaia* DR3 are shown as orange circles. Blue contours trace the density of the *Gaia* subset, and red contours trace the density of the non-*Gaia* subset. The marginal histograms for z_{vi} (above) and H_E (right) show the full sample (black steps), the *Gaia* subset (blue steps), and the non-*Gaia* subset (orange steps). *Bottom*: One-dimensional histograms of I_E , Y_E , and J_E for the same three samples. All histograms are normalised to unit area.

a percentile-based scaling,

$$F_{\text{norm}}(\lambda) = \frac{F_{\text{obs}}(\lambda) - P_{25}(F_{\text{obs}}(\lambda))}{P_{95}(F_{\text{obs}}(\lambda)) - P_{25}(F_{\text{obs}}(\lambda))}, \quad (7)$$

where $F_{\text{obs}}(\lambda)$ is the original flux density array, and P_{25} and P_{95} denote the 25th and 95th percentiles of $F_{\text{obs}}(\lambda)$, respectively. This normalisation reduces the continuum level and enhances the emission lines on the stacked plot.

From Fig. 7, the most prominent emission line among all spectra is $\text{H}\alpha$, which lies in the wavelength range of $[12047, 18734] \text{ \AA}$ at redshift $0.83 \leq z \leq 1.85$. Other prominent emission lines include $\text{He I} + \text{Pa } \gamma$ at $0.11 < z < 0.73$, $\text{H}\beta + [\text{O III}]$ at $1.5 \leq z \leq 2.8$, and Mg II at $z > 3.3$. While present in the emission-line map, $\text{H}\gamma$ becomes faint at high redshift, especially when $\text{H}\beta + [\text{O III}]$ move out of the observed wavelength range. The low equivalent width (EW) and S/N of $\text{H}\gamma$ pose a challenge to the redshift determination at $2.8 < z < 3.3$, which can only be alleviated by combining with data at other wavelengths. In future data releases of *Euclid*, the blue-grism data of NISP, covering 926–1366 nm, will be available for *Euclid* Deep and Auxiliary fields (Euclid Collaboration: Mellier et al. 2025). Inclusion of the blue-grism data will increase the efficiency of source identification and redshift determination in these deep fields.

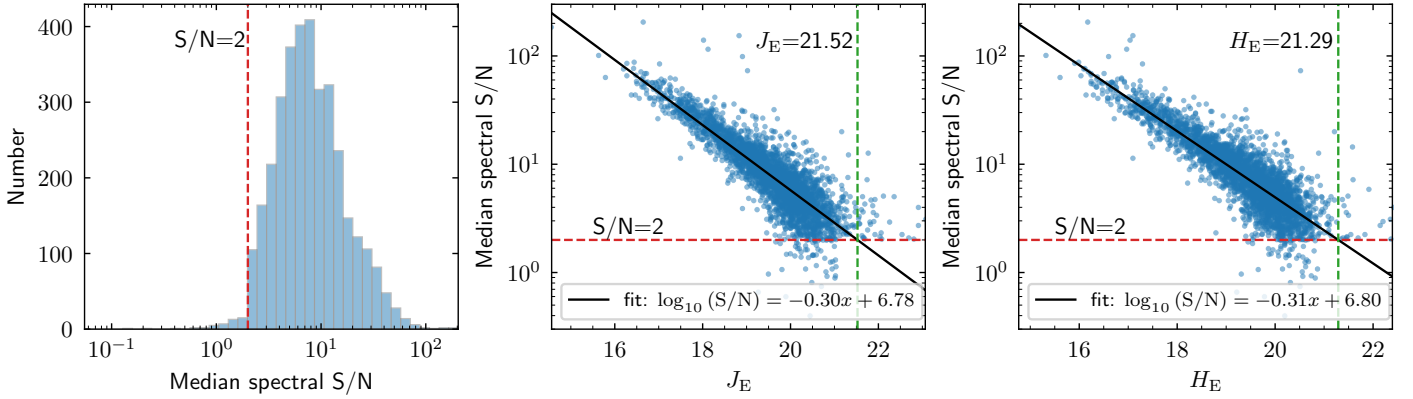


Fig. 4. *Left:* Distribution of the median spectral S/N within [12047, 18734] Å for visually identified quasars. The dashed red line at S/N = 2 marks an empirical threshold below which the number of successful identifications declines rapidly. *Middle and right:* Median S/N versus J_E and H_E magnitudes, respectively. Blue points show individual quasars, and the black lines indicate linear fits to $\log_{10}(S/N)$, with the fitted relations annotated. The dashed red line again marks S/N = 2, and its intersection with the fit defines empirical limiting magnitudes of $J_E \approx 21.5$ and $H_E \approx 21.3$, shown by dashed green lines. These values represent practical limits for reliable spectral identification of quasars in Q1.

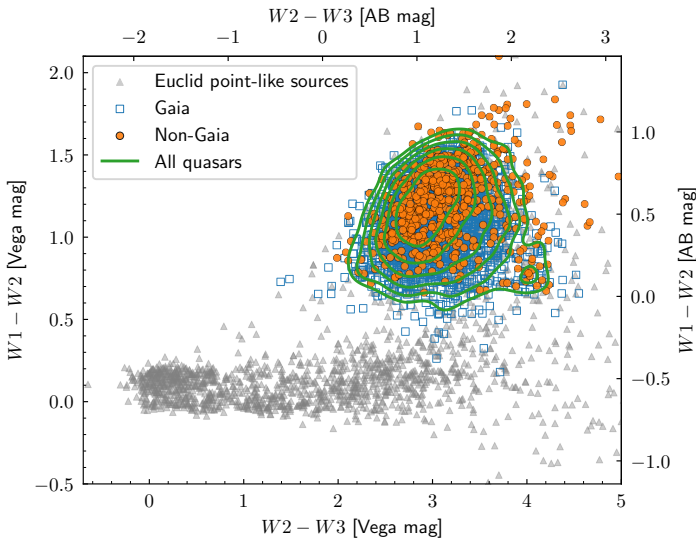


Fig. 5. $W1 - W2$ versus $W2 - W3$ colour-colour diagram of spectroscopically identified quasars in this work, where *Gaia*-detected sources are shown as blue open squares, and sources not in *Gaia* DR3 are shown as orange circles. The density distribution of the full identified quasar sample is indicated with green contour lines. *Euclid* point-like sources are shown as grey triangles. To ensure the reliability of the colours, only sources with adequate S/N ($w1snr > 5$, $w2snr > 5$, and $w3snr > 3$) are shown.

4.3. Composite Euclid spectra of the golden sample of bright quasars

A mean or median composite spectrum of a sample of quasars is useful for understanding the average spectral properties across a wide wavelength range, and for providing a representative spectrum template (e.g. Vanden Berk et al. 2001; Glikman et al. 2006). To generate the mean and median composite quasar spectra spanning the rest-frame optical to NIR wavelengths, we adopted a procedure similar to that described by Vanden Berk et al. (2001), with a 1D ‘drizzle’ technique to improve the sampling of the low-resolution spectra. We briefly introduce the procedure below, and refer to Appendix C for technical details. We first corrected the Milky Way dust extinction of the spectra using $E(B-V)$ values from the Galactic dust map produced by Planck Collaboration XLVIII (2016), and the extinc-

tion law of Gordon et al. (2023) assuming $R_V = 3.1$. The packages *dustmaps* (Green 2018) and *dust_extinction* (Gordon 2024) are used for the correction. The spectra are then sorted in ascending order of redshift. The spectrum with the lowest redshift is normalised arbitrarily to have a unit mean flux density. Each subsequent spectrum is shifted to the rest-frame based on its measured redshift and normalised to match the mean flux density of the mean composite built from all lower-redshift spectra, within their overlapping wavelength range.

All spectra were resampled onto a common rest-frame wavelength grid with a constant bin size of $\Delta\lambda = 4$ Å, using a flux-conserving re-binning method to preserve the integrated flux density in each bin. This process is mathematically equivalent to a 1D version of the Drizzle algorithm (Fruchter & Hook 2002), which has been widely applied to image reconstructions by combining dithered, undersampled images in HST and JWST surveys (e.g. Koekemoer et al. 2011; Williams et al. 2023; Bagley et al. 2023). We set $\Delta\lambda = 4$ Å to match the smallest native rest-frame pixel size near the blue end while maintaining higher per-pixel S/N at the highest redshifts. For reference, the observed sampling of 13.4 Å at $z = 2.35$ corresponds to $13.4 \text{ Å} / (1 + 2.35) = 4 \text{ Å}$ in the rest frame, and the blue-end wavelength at this redshift is $12047 \text{ Å} / (1 + 2.35) = 3596 \text{ Å}$. Consequently, the common 4-Å grid oversamples the data at longer rest wavelengths and applies mild downsampling at the shortest rest wavelengths contributed by the highest-redshift quasars. By stacking the oversampled rest-frame spectra of different redshifts through drizzle, the composite recovers some information that is lost in the individual spectra at rest-frame $\lambda > 3596$ Å, giving finer emission line features than individual ones.

The arithmetic mean composite (mean composite for short) is calculated as the equal-weight mean of the sigma-clipped normalised flux densities in each bin. The uncertainties of the mean composite are computed through error propagation, assuming uncorrelated input pixels. In parallel, the root mean square (RMS) flux that quantifies the object-to-object dispersion is recorded. A median and a geometric mean composite spectrum are also generated with the same sigma-clipped data as used for the mean composite. The mean, median, and geometric mean composite spectra, along with the uncertainties of the arithmetic and geometric mean composites, RMS flux, S/N, and number of spectra in each wavelength bin (N_{spec}), are tabulated in Table 2.

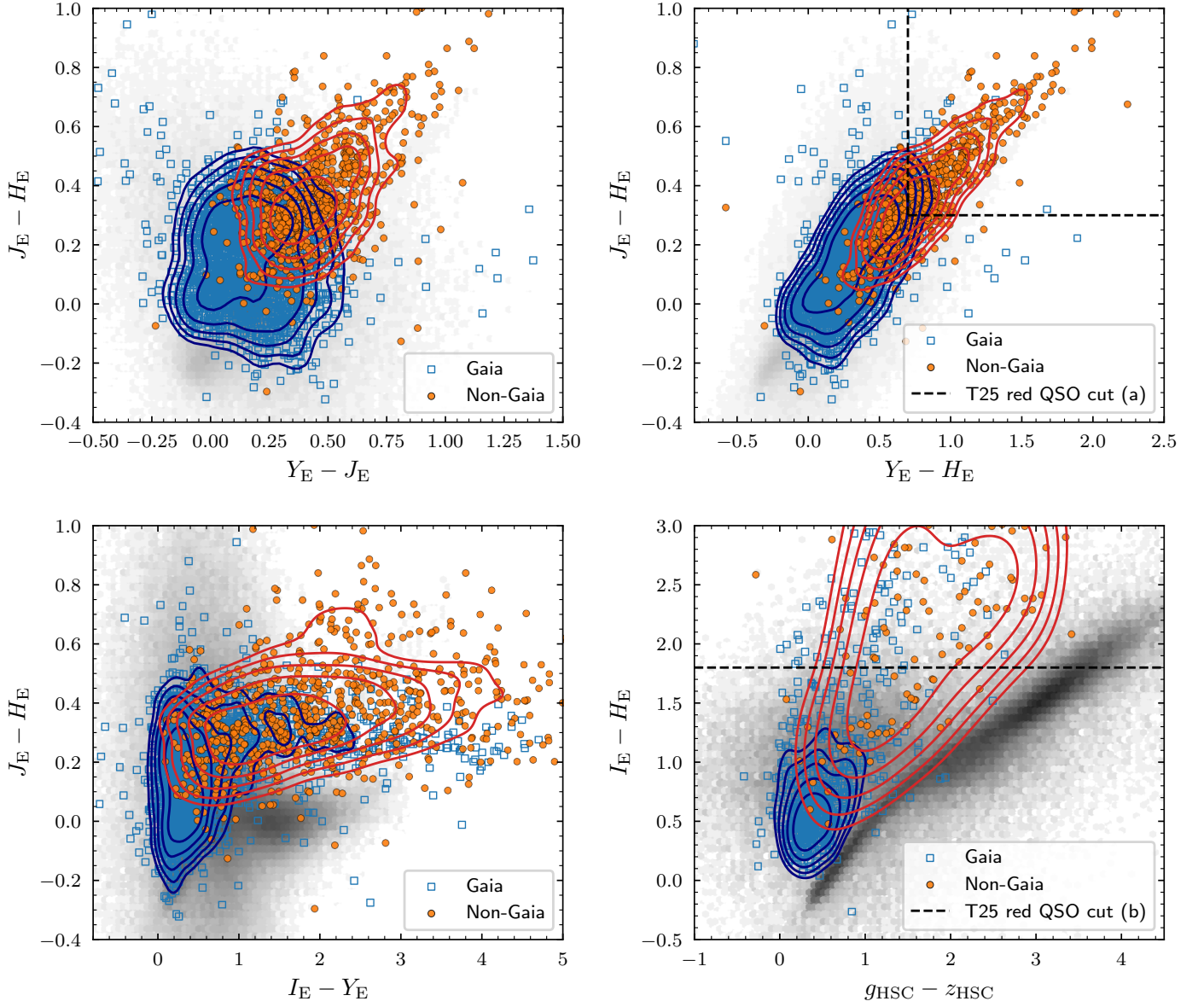


Fig. 6. *Euclid* and external (Hyper Suprime-Cam, HSC; Miyazaki et al. 2018; Aihara et al. 2018) colour-colour diagrams of spectroscopically identified quasars in this work, where *Gaia*-detected sources are shown as blue open squares, and sources not in *Gaia* DR3 are shown as orange circles. The density distribution of the *Gaia*-detected sample is indicated with blue contour lines, and that of the non-*Gaia* sample is indicated with red contour lines. The density of *Euclid* point-like sources is shown as grey hexagonal binning plots in the background.

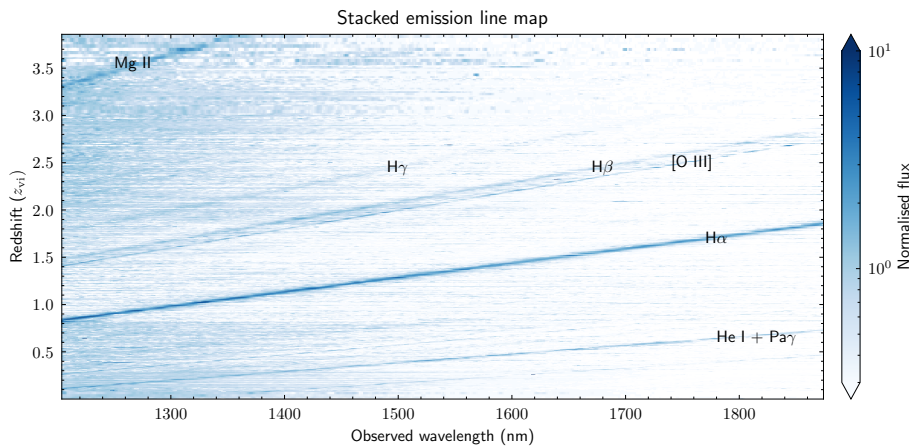


Fig. 7. Stacked emission-line map of the golden sample of 2868 visually identified quasars. Each row of the pixels represents a spectrum normalised using its 25th and 95th percentiles. Prominent emission lines are marked with text labels.

Table 2. Mean, geometric mean, and median *Euclid* Q1 quasar composite spectra, along with RMS, S/N and the number of spectra in each wavelength bin (N_{spec}).

Wavelength (nm)	Mean F_λ (arb.)	Geometric mean F_λ (arb.)	Median F_λ (arb.)	RMS (arb.)	S/N	N_{spec}
240.0	15.00 ± 1.08	14.71 ± 1.13	14.35	2.97	13.83	5
240.4	13.11 ± 1.03	12.51 ± 0.96	13.44	3.70	12.74	5
240.8	15.90 ± 1.22	15.65 ± 1.18	16.19	2.83	12.99	5
241.2	14.82 ± 1.16	14.13 ± 1.26	14.39	4.59	12.73	5
241.6	14.93 ± 1.06	14.87 ± 1.05	14.59	1.35	14.12	5
...	...	...	...	...	...	...
1698.4	0.83 ± 0.01	0.83 ± 0.02	0.84	0.06	55.20	9
1698.8	0.83 ± 0.02	0.82 ± 0.02	0.83	0.06	50.31	9
1699.2	0.77 ± 0.02	0.72 ± 0.02	0.83	0.20	46.35	9
1699.6	0.84 ± 0.02	0.84 ± 0.02	0.83	0.05	48.51	8
1700.0	0.86 ± 0.02	0.86 ± 0.02	0.86	0.08	50.59	8

Notes. Table 2 is available in its entirety at the CDS. A portion is shown here for guidance regarding its form and content.

Figure 8 presents the mean and median composite spectra derived from our sample, together with the RMS scatter around the mean composite. The mean quasar composite spectrum from Glikman et al. (2006), the mean composite of type 1 AGNs from Euclid Collaboration: Lusso et al. (2024), and a mean composite spectrum constructed by Euclid Collaboration: Lusso et al. (2024) using datasets from Landt et al. (2008, 2011, 2013), are also shown for comparison. Prominent emission lines are marked in the composite spectra for reference, including Mg II λ 2800, [O II] λ 3728, H γ , H β , [O III] λ 4960, 5008, H α , He I, Pa γ , and Pa β . The mean and median composite spectra are consistent with each other.

The number of contributing spectra (N_{spec}), the S/N, and the RMS flux of the mean composite as functions of rest-frame wavelength are shown in Fig. 9. Between 100 and more than 1000 spectra contribute to each wavelength bin over the range 0.32–1.48 μm , which yields S/N values above 100 in the majority of bins. The RMS about the mean is typically between 0.1 and 1 in the continuum (in units of the mean flux), and usually increases by less than an order of magnitude at the positions of emission lines, except around Pa β where the RMS rises by about an order of magnitude. This behaviour indicates a moderate level of intrinsic quasar diversity that dominates the variance near strong lines.

The Glikman et al. (2006) composite (up to 3.52 μm) is constructed using data from 27 bright quasars observed with the NASA IRTF, and the Euclid Collaboration: Lusso et al. (2024) type 1 composite (up to 3.61 μm) is built with 23 quasars from Glikman et al. (2006) and nine additional hard-X-ray-selected AGNs observed with the folded-port infrared echellette (FIRE; Simcoe et al. 2008) from Ricci et al. (2022). The Landt et al. (2008, 2011, 2013) composite covering 0.75–2.3 μm is based on 29 well-known local type 1 AGN observed at the NASA IRTF and the Gemini North observatory. At $\lambda < 0.7 \mu\text{m}$, the mean composite spectrum from this work (hereafter the *Euclid* composite) shows a consistent continuum slope to the composites from Glikman et al. (2006) and Euclid Collaboration: Lusso et al. (2024). At longer wavelengths, both Glikman et al. (2006) and Euclid Collaboration: Lusso et al. (2024) composites display residual telluric absorption features, while the *Euclid* composite shows a clean continuum, free of telluric absorption. The

Landt et al. (2008, 2011, 2013) composite shows a smoother continuum than the other two ground-based composite spectra, and a steeper slope at $0.75 < \lambda < 0.98 \mu\text{m}$.

To quantitatively compare the NIR continuum shape with previous work, we fitted a broken power law in F_λ to the *Euclid* Q1 geometric mean composite and to three published quasar composites over the wavelength range 0.75–1.35 μm (Fig. 10), using the spectral fitting package QSOFITMORE (Fu 2025). The break was fixed at 0.98 μm based on visual inspection. The upper limit of 1.35 μm was chosen to exclude the strong long-wavelength upturn of the Euclid Collaboration: Lusso et al. (2024) composite and to reduce the impact of small-number statistics in the Q1 composite. We parameterised the power-law model as $F_\lambda \propto \lambda^{\alpha_\lambda}$ and describe the broken power law with indices $\alpha_{\lambda,1}$ and $\alpha_{\lambda,2}$ blueward and redward of the break, respectively.

For the *Euclid* Q1 composite, we obtain $\alpha_{\lambda,1} = -0.97$ and $\alpha_{\lambda,2} = -0.73$. The type 1 composite from Euclid Collaboration: Lusso et al. (2024) has similar slopes, $\alpha_{\lambda,1} = -1.00$ and $\alpha_{\lambda,2} = -0.57$, while the Glikman et al. (2006) geometric mean composite yields $\alpha_{\lambda,1} = -1.25$ and $\alpha_{\lambda,2} = -0.68$. The mean Landt et al. composite shows a rather steep blue segment with $\alpha_{\lambda,1} = -2.28$ over 0.75–0.98 μm , then turns over to a much flatter slope, $\alpha_{\lambda,2} = -0.56$, over 0.98–1.35 μm . In the F_λ representation, all four composites show a mild flattening of the NIR continuum toward longer wavelengths. For comparison with the literature, the corresponding α_ν values (computed via $\alpha_\nu = -[\alpha_\lambda + 2]$) are listed in parentheses in Fig. 10. Our *Euclid* composite has the flattest continuum in the NIR (with the smallest change of slopes at the break), while the Landt et al. mean composite exhibits the largest curvature (change of slopes). As shown by Landt et al. (2013), the NIR spectral energy distribution (SED) of AGNs affected by strong host galaxy light are much flatter than those of AGNs with low host contribution. The flatness of the NIR continuum of our *Euclid* composite is therefore most likely due to the host galaxy light contribution of the low-redshift AGNs.

The observed central wavelengths of selected emission lines (Table 3) were measured from the mean composite with spectral fitting using QSOFITMORE (Fu 2025). For broad emission lines, including Mg II, H β , H α , O I, He I, Pa δ , Pa γ , and Pa β , we fitted a narrow Gaussian component ($\text{FWHM} \leq 1200 \text{ km s}^{-1}$) and 2–3 broad Gaussian components ($\text{FWHM} > 1200 \text{ km s}^{-1}$) to each line profile, and report the central wavelengths of the narrow components. For narrow (forbidden) lines, including [O II], [Ne III], [O III], and [S III], we fitted only one narrow component to each profile. The uncertainties are the standard deviations of a Monte Carlo simulation of 50 fits. As shown in Table 3, the line centres of narrow lines typically show lower uncertainties than those of the broad lines ([O III] λ 5008 has the lowest wavelength uncertainty), indicating the significance of narrow lines in redshift determination.

5. Discussion

5.1. Morphological properties of the bright quasar sample

With a pixel scale of 0".1 and stable PSF of VIS imaging, *Euclid* resolves obvious galaxy structure from the local Universe to at least $z = 1.5$ (with particularly rich statistics at $0.3 < z < 0.7$; see e.g. Euclid Collaboration: Walmsley et al. 2026). These data enable the detailed structural characterisation of quasar hosts through uniform measurements of Sérsic-based and model-independent parameters (Euclid Collaboration: Quilley et al.

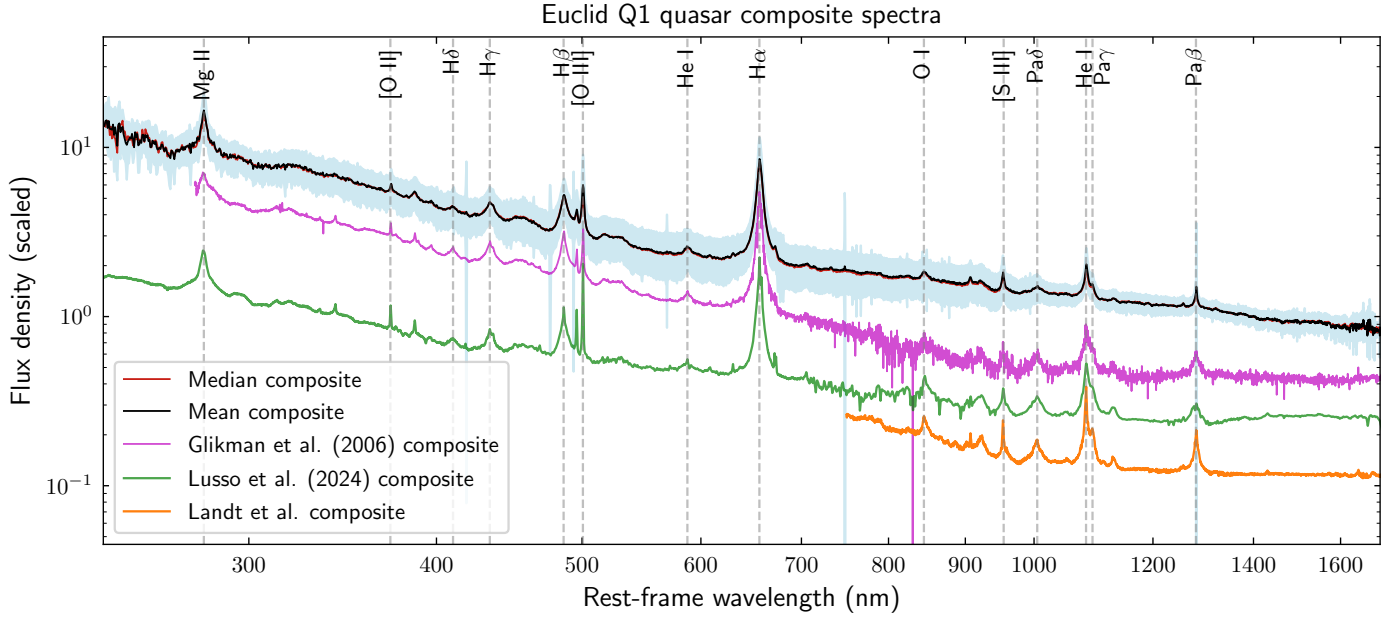


Fig. 8. Mean and median composite quasar spectra constructed from our golden sample (black and red curves, respectively), with the RMS scatter around the mean composite indicated by the shaded blue region. The quasar composite spectrum from Glikman et al. (2006) is plotted in magenta, the type 1 AGN composite from Euclid Collaboration: Lusso et al. (2024) is plotted in green, and the mean composite constructed using data from Landt et al. (2008, 2011, 2013) is plotted in orange. Prominent emission lines are marked for reference.

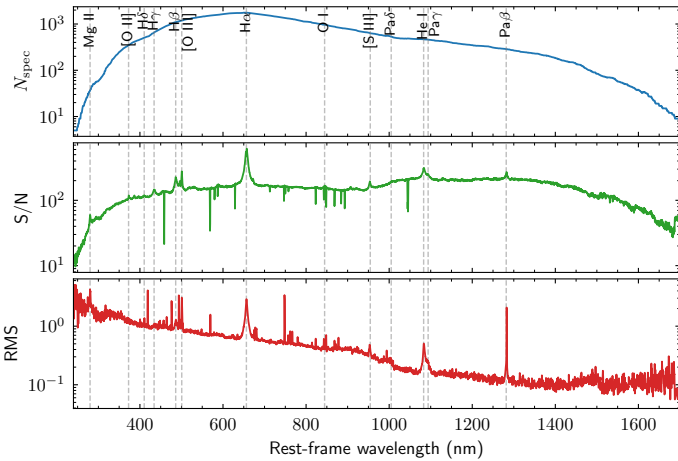


Fig. 9. Diagnostics of the Euclid Q1 mean quasar composite as a function of rest-frame wavelength for a bin size of $\Delta\lambda = 4 \text{ \AA}$. *Top:* number of spectra contributing to each wavelength bin. *Middle:* S/N of the mean composite per bin. *Bottom:* RMS dispersion of the contributing spectra about the mean. Vertical dashed lines mark the rest wavelengths of prominent emission lines, which are labelled in the top panel.

2026; Euclid Collaboration: Romelli et al. 2026), as well as the separation of compact nuclear light from extended galaxy emission (Euclid Collaboration: Margalef-Bentabol et al. 2026).

To take advantage of the detailed structural information provided by the VIS images, we examined two redshift regimes separately: (i) a low-redshift ($z < 0.5$) subset with 341 sources, and (ii) an intermediate-redshift ($0.5 < z < 2$) subset with 2361 sources. Both subsets are selected to have valid morphological parameters measured from VIS imaging, including parameters from the single-component Sérsic fit (Euclid Collaboration: Quilley et al. 2026): Sérsic index (n_{VIS}), effective radius ($R_{e,\text{VIS}}$), and axis ratio (ρ_{VIS}); and the model-independent CAS parameters (Conselice

2003; Euclid Collaboration: Romelli et al. 2026): concentration, asymmetry, and clumpiness (smoothness). The intermediate-redshift subset is further supplemented with a VIS PSF fraction (f_{PSF}), derived by the deep-learning-based f_{PSF} prediction model trained on simulated galaxy images with different levels of f_{PSF} added to them (Euclid Collaboration: Margalef-Bentabol et al. 2026). We choose the redshift range $0.5 < z < 2$ because the model was trained and validated in this range.

As shown in Fig. 11 and Table 4, the low-redshift population shows a large source extent with median $\mu_{\text{max}} - \text{mag} = -1.67$, significantly higher than the median $\mu_{\text{max}} - \text{mag} = -2.87$ of the intermediate-redshift sample. The $\mu_{\text{max}} - \text{mag}$ value anticorrelates with the Sérsic index and concentration. In particular, the most compact sources in this population ($\mu_{\text{max}} - \text{mag} \approx -2.5$) have a Sérsic index around 5.5, which is the upper limit set by Euclid Collaboration: Quilley et al. (2026). They have suggested that fits with Sérsic indices above 5.45 should be removed from any Sérsic-based analysis. In this low-redshift population, 50% of the sources have n_{VIS} higher than 5.45. Such near-boundary Sérsic indices indicate that the single-component Sérsic model cannot describe the light profiles of AGNs with bright cores. When requiring $n_{\text{VIS}} < 5.45$, the median n_{VIS} is 2.31, which is still significantly higher than the peak value of 0.8 among all galaxies in Euclid Collaboration: Quilley et al. (2026). The median $R_{e,\text{VIS}}$ of $n_{\text{VIS}} < 5.45$ sources is $1''.0$, identical to the peak value of $R_{e,\text{VIS}}$ in Euclid Collaboration: Quilley et al. (2026).

The CAS parameters further characterise the resolved structure of the low-redshift quasars. The median concentration ($C = 4.27$) is high compared to the general galaxy population (e.g. $C \approx 2.5$, see Euclid Collaboration: Quilley et al. 2026) but with significant scatter, reflecting the coexistence of bright nuclear light and extended hosts. Asymmetry values are moderate (median $A = 0.36$) with a tail to $A > 1$, indicating disturbed morphologies or nearby companions. The clumpiness parameter is elevated (median $S = 0.19$), suggesting clumpy emission in many cases, possibly due to ongoing star formation (e.g. Conselice 2003, 2014). Overall, at $z < 0.5$, the host galaxies of

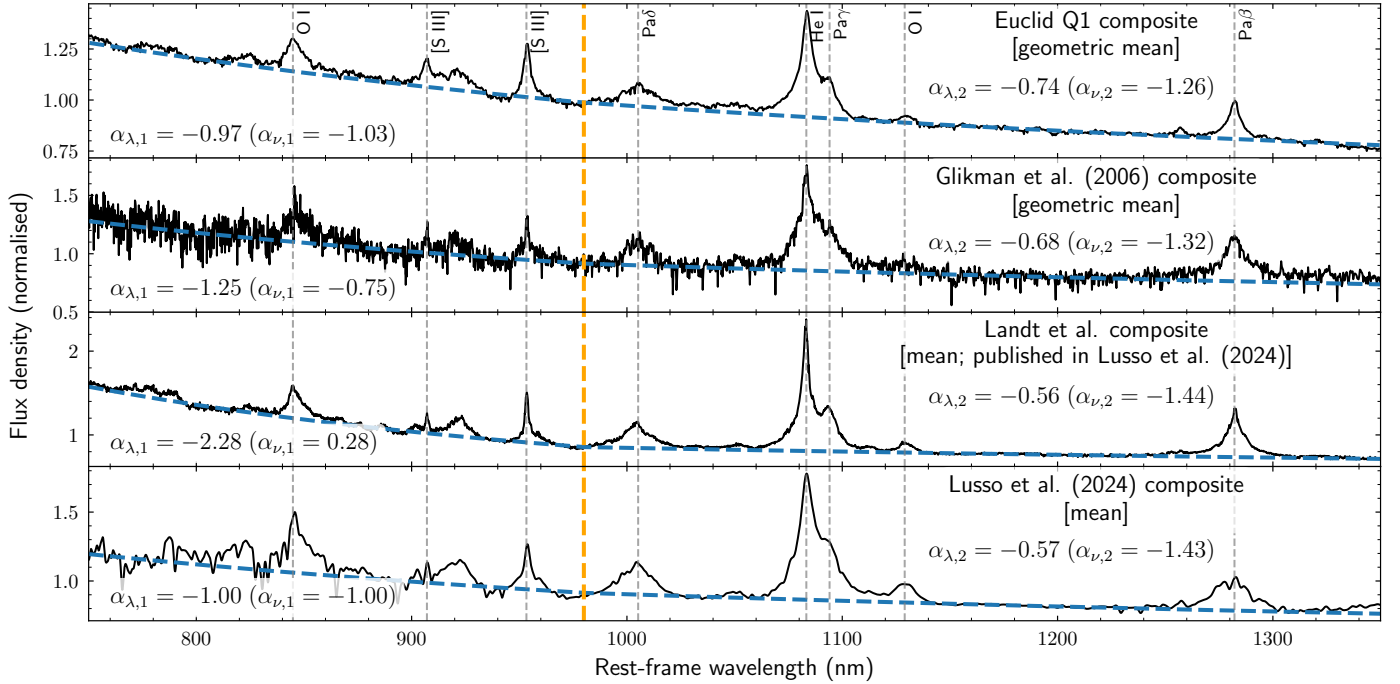


Fig. 10. *Euclid* Q1 geometric mean composite and three literature quasar composites, with NIR continua fitted with broken power laws over 0.75–1.35 μm . From top to bottom, the panels show the *Euclid* Q1 composite (geometric mean), the geometric mean composite from Glikman et al. (2006), the mean Landt et al. composite as published by Euclid Collaboration: Lusso et al. (2024), and the type 1 AGN composite from Euclid Collaboration: Lusso et al. (2024). The black curves show the composite spectra, the dashed blue lines show the best-fitting broken power laws in F_λ , and the vertical dashed orange lines indicate the break wavelength of 980 nm. The annotated indices $\alpha_{\lambda,1}$ ($\alpha_{\nu,1}$) and $\alpha_{\lambda,2}$ ($\alpha_{\nu,2}$) are the corresponding spectral slopes in F_λ (F_ν) blueward and redward of the break. Selected emission lines are labelled in the top panel for reference.

Table 3. Major emission lines identified from the mean composite quasar spectrum.

Line ID	λ_{vac} (Å)	λ_{obs} (Å)	Intensity [100F/H α]	EW (Å)	FWHM (km s $^{-1}$)
Mg II	2798.75	2799.37 $\pm$ 0.8	16.0 $\pm$ 8.7	7.5	6444.9
[O II]	3728.48	3731.54 $\pm$ 0.45	2.2 $\pm$ 0.2	1.9	1200.0
[Ne III]	3868.58	3868.73 $\pm$ 0.85	2.4 $\pm$ 0.3	2.3	1200.0
H γ	4341.68	4338.57 $\pm$ 1.33	10.3 $\pm$ 1.0	12.5	4195.5
H β	4862.68	4863.73 $\pm$ 0.86	38.6 $\pm$ 3.9	61.3	8241.9
[O III]	4960.3	4961.25 $\pm$ 0.3	4.5 $\pm$ 0.6	7.2	1200.0
[O III]	5008.24	5007.79 $\pm$ 0.07	13.6 $\pm$ 1.4	21.7	1200.0
He I	5877.29	5873.04 $\pm$ 5.88	0.1 $\pm$ 0.1	0.2	316.2
H α	6564.61	6564.63 $\pm$ 1.34	100.0	219.6	5871.2
O I	8448.8	8445.11 $\pm$ 3.37	1.0 $\pm$ 1.0	2.9	3209.3
[S III]	9071.1	9068.76 $\pm$ 0.66	1.2 $\pm$ 0.1	3.7	1200.0
[S III]	9533.2	9535.53 $\pm$ 0.35	3.0 $\pm$ 0.3	9.7	1185.9
Pa δ	10052.1	10055.74 $\pm$ 1.55	4.4 $\pm$ 0.5	15.1	5378.1
He I	10833.2	10836.36 $\pm$ 0.21	3.0 $\pm$ 0.3	10.9	1016.5
Pa γ	10941.1	10941.91 $\pm$ 0.75	0.6 $\pm$ 0.2	2.3	972.8
O I	11290.0	11297.54 $\pm$ 2.38	0.4 $\pm$ 0.1	1.7	2825.8
Pa β	12821.6	12825.47 $\pm$ 0.23	5.2 $\pm$ 0.5	22.3	2825.8

Notes. For each line, its vacuum wavelength (λ_{vac}) is listed, along with observed wavelength (λ_{obs}) with 1σ uncertainty, relative intensity to H α , EW, and full width at half maximum (FWHM).

bright quasars are frequently resolved and often display interacting features and star-forming clumps.

Compared to the low-redshift subset, the intermediate-redshift quasars at $0.5 < z < 2$ are markedly more compact (see Fig. 12 and Table 4). More than 90% of the sources (2031) have near-boundary Sérsic indices of $n_{\text{VIS}} > 5.45$, which results

in a median $n_{\text{VIS}} = 5.5$ and a median $R_{\text{e,VIS}} = 0''.01$ (1/10 of the VIS pixel size). Such a high fraction of saturated Sérsic fits indicates that the single-component Sérsic model cannot capture the unresolved nuclear dominance of these sources. When restricting to $n_{\text{VIS}} < 5.45$, the median Sérsic index decreases to 2.25 and the median effective radius increases to $0''.42$, values more consistent with resolved galaxy light profiles.

The VIS PSF fraction (f_{PSF}) provides a more direct handle on this unresolved light. Its distribution peaks around 0.8 with another weaker peak around 0.1 (Fig. 12), confirming that most sources are dominated by their cores, consistent with their Sérsic fits converging to extreme values. Sources with high f_{PSF} values also correspond to the most point-like objects, characterised by $\mu_{\text{max}} - \text{mag} \approx -2.9$. However, low f_{PSF} values (e.g. $f_{\text{PSF}} < 0.1$) do not necessarily imply the absence of AGN activity, because uncertainties of the measurements could dominate at such low levels.

As shown in Fig. 12 and Table 4, the CAS parameters of the intermediate-redshift quasars show very different distributions compared to those of the low-redshift quasars. Highly left-skewed distributions of concentration (median $C = 2.56$) and clumpiness (median $S = 0.08$) are seen in intermediate-redshift quasars, in contrast to the nearly symmetric distributions of the two parameters for the low-redshift quasars (median $C = 4.27$ and median $S = 0.19$). The concentration and clumpiness distributions of the intermediate-redshift quasars appear unphysical, because quasar hosts are expected to be more concentrated and clumpier towards higher redshift. This discrepancy likely arises because the nuclear PSF components were not removed before the measurements. Indeed, sources with low nuclear fractions ($f_{\text{PSF}} \lesssim 0.3$) show more realistic Sérsic indices and CAS

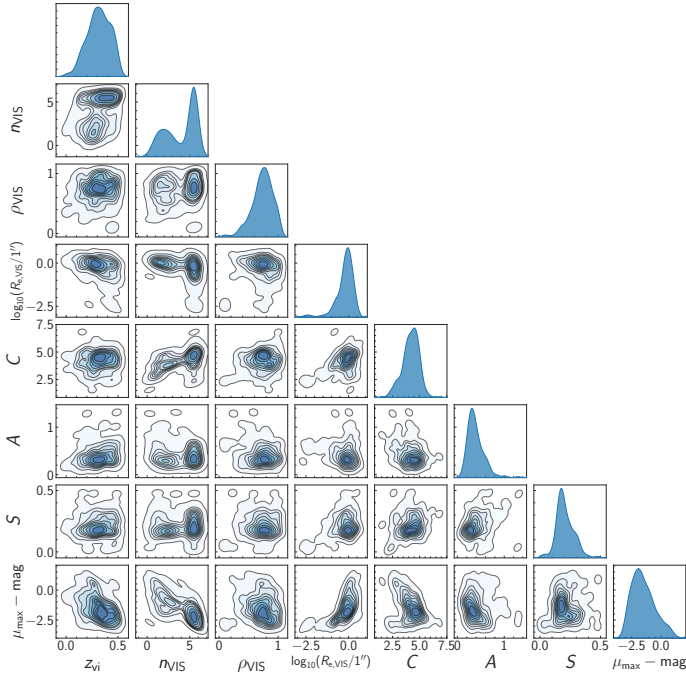


Fig. 11. Corner plot showing the joint distributions of morphological parameters measured from *Euclid* VIS imaging for 341 sources at $z < 0.5$. Displayed parameters include visual-inspection redshift (z_{vl}), Sérsic index (n_{VIS}), axis ratio (ρ_{VIS}), logarithmic effective radius ($\log_{10}(R_{\text{e,VIS}}/1'')$), concentration (C), asymmetry (A), clumpiness (S), and $\mu_{\text{max}} - \text{mag}$.

Table 4. Median VIS morphology parameters and f_{PSF} .

Parameter	$z < 0.5$		$0.5 < z < 2$	
	all	$n_{\text{VIS}} < 5.45$	all	$n_{\text{VIS}} < 5.45$
N	341	171	2361	230
n_{VIS}	5.35	2.31	5.50	2.25
ρ_{VIS}	0.76	0.71	0.69	0.67
$R_{\text{e,VIS}}$	0.78	1.00	0.01	0.42
C	4.27	3.96	2.56	2.93
A	0.36	0.33	0.55	0.50
S	0.19	0.19	0.08	0.19
$\mu_{\text{max}} - \text{mag}$	-1.67	-0.94	-2.87	-1.68
f_{PSF}	–	–	0.76	0.16

Notes. f_{PSF} is available only for $0.5 < z < 2$.

parameters, reinforcing that uncorrected nuclear light biases the structural measurements of the host galaxies.

Taken together, these results demonstrate that VIS morphology captures a strong redshift dependence in the apparent structure of bright quasars: (i) at $z < 0.5$, obvious host galaxy structures are frequently resolved, and a single Sérsic model can describe the light profiles of half of the sources; (ii) at $0.5 < z < 2$, the nuclear component dominates, driving Sérsic indices of 90% of the sources to their upper limits. In this regime ($0.5 < z < 2$), f_{PSF} offers a more reliable compactness measure and light profile characterisation than the single Sérsic fits, allowing us to quantify the balance between nuclear and host emission across the bright quasar sample. Representative *Euclid* 1D spectra and VIS cutouts illustrating these trends are shown in Appendix D for low-redshift sources ($z < 0.5$; Fig. D.1) and for $0.5 < z < 2$ quasars spanning different f_{PSF} levels (Fig. D.2). A

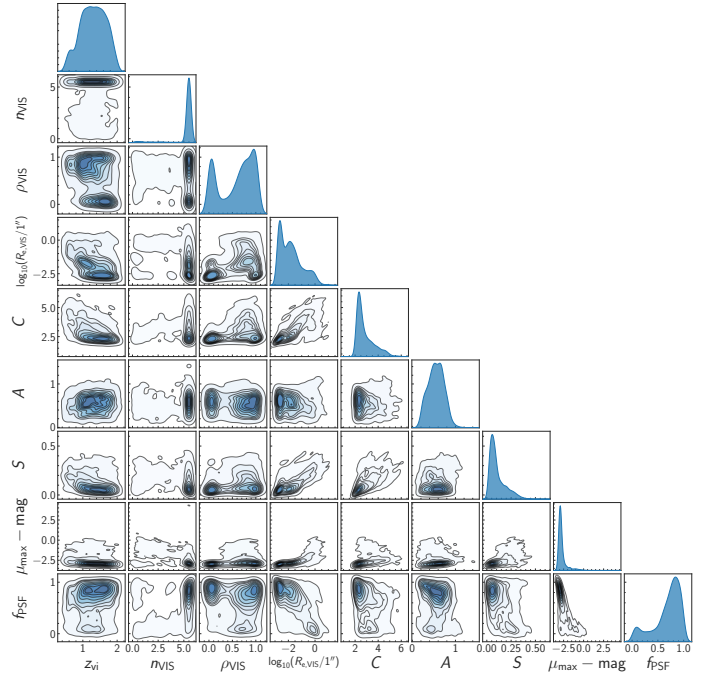


Fig. 12. Same as Fig. 11, but for 2361 sources at $0.5 < z < 2$ with measurements of AGN PSF contribution fraction (f_{PSF}).

more detailed analysis of AGN host galaxy morphology, including re-measurement of morphological parameters with methods tailored to AGN host galaxies, and dedicated tests to quantify measurement sensitivity as a function of host galaxy surface brightness, redshift, and nucleus-to-host contrast, will be presented in a follow-up study.

5.2. *Euclid* colour space, multi-wavelength selection, and future improvements

The advent of *Euclid*'s high-quality, wide-area near-infrared photometry opens a novel colour space for quasar selection, providing new opportunities to discover dust-reddened or intrinsically red quasars that are missed by traditional optical or ultraviolet (UV) selection. The extended wavelength coverage of the *Euclid* NIR bands, combined with precise VIS photometry, enables more robust discrimination of quasars from stars and galaxies, particularly at redshifts and extinction levels where classical techniques are less effective. Our initial exploration demonstrates that red quasars occupy distinct loci in *Euclid* colour-colour diagrams, highlighting the potential of the survey to identify populations that are heavily underrepresented in previous optical surveys.

Mid-infrared photometry from WISE has already proven highly effective for selecting luminous AGNs and quasars via their characteristic red MIR colours. Our analysis confirms that WISE-based colour cuts efficiently select AGN-dominated sources, as reflected in the properties of our pre-selected sample. However, WISE alone cannot distinguish between the most dust-obscured and less obscured quasars. The synergy between *Euclid* and WISE thus offers a powerful selection strategy: by combining WISE MIR colours with deep VIS and NIR photometry from *Euclid*, future quasar searches can achieve high completeness, efficiently probing obscured AGN populations.

Moreover, incorporating multi-band photometric data into the spectroscopic identification process will substantially

improve the completeness and reliability of the quasar sample. Photometric pre-selection not only increases the likelihood of identifying genuine quasars, especially those with atypical SEDs, but also enhances redshift determination by providing prior constraints and mitigating degeneracies in the spectral fitting. The inclusion of *Euclid* and external photometry will be essential for robust identification and breaking redshift degeneracies from only one emission line.

A limitation of the present work is that the initial sample relies on pre-selection using *Gaia* and WISE catalogues, which may miss quasars outside the colour selection windows or at the limits of the *Gaia* and WISE sensitivity. In future *Euclid* data releases, the construction of a more complete quasar sample will be enabled by combining traditional pre-selected candidates with new candidates identified from *Euclid* photometry and by the SPE pipeline. This composite approach, leveraging the full capability of the *Euclid* data, will allow for the recovery of quasars with a wider range of colours, luminosities, and host properties.

In addition, we note that measuring the intrinsic spectral properties of quasars requires careful treatment of the host galaxy contribution. As demonstrated in Sect. 5.1, even in AGN-dominated sources, the host can provide a significant fraction of the observed flux, especially in the VIS band. Correctly separating nuclear and host emission is challenging and may introduce additional uncertainties in line and continuum measurements. Future analyses will benefit from improved morphological decomposition and SED modelling to robustly isolate the quasar component and fully exploit *Euclid*'s spectrophotometric capabilities.

6. Conclusions

In this work, we have presented a large sample of 3468 bright quasars covering the redshift range of $0 < z \leq 4.8$ identified with red-grism spectroscopy of the first *Euclid* Quick Data Release, including 2686 sources with new spectroscopic identifications relative to existing public compilations. To ensure identification efficiency, we focused on a high-purity input quasar candidate sample of 9214 sources based on GDR3-QSOs (Quaia, CatNorth, and CatSouth) and the AllWISE R90 AGN table. Through a template matching process based on the Pearson correlation function and a visual inspection campaign, we labelled quasars and determine their spectroscopic redshifts. The success rate of the spectroscopic identification is 38%.

From the identified sample, we estimated an empirical spectroscopic depth of $J_E \lesssim 21.5$ and $H_E \lesssim 21.3$ at the sensitivity of the Wide Field Survey, beyond which the number of securely identified quasars declines sharply. Our investigation of the novel *Euclid* colour space demonstrates its power in uncovering redder, dust-obscured quasars that may be missed by traditional optical and MIR selections. The synergy between *Euclid* and WISE photometry promises even greater completeness and diversity in future quasar surveys.

We constructed the first *Euclid* composite spectrum of bright quasars, covering rest-frame NUV to NIR wavelengths and free from telluric absorption, providing a valuable benchmark for future spectral studies. The spectroscopic bright quasar catalogue of this work, and the composite quasar spectrum, will be available at CDS. In addition, the spectral properties of the quasars will be derived by multi-component fitting (Euclid Collaboration: Calhau et al. in preparation).

We characterised VIS morphologies using single-Sérsic and model-independent (CAS) parameters, supplemented by a deep-learning PSF fraction, f_{PSF} . At low redshift ($z < 0.5$), obvious host structures are frequently resolved, and a single Sérsic model

describes the light profiles of about half of the sources. At intermediate redshift ($0.5 < z < 2$), the nuclear component dominates, driving Sérsic indices of roughly 90% of the sources to the upper bound; in this regime f_{PSF} provides a more reliable compactness measure and is used to quantify the balance between nuclear and host emission across the sample.

Finally, we discuss current limitations, including the pre-selection bias and the challenges of host-nucleus decomposition, and outline pathways for improvement with upcoming *Euclid* data releases and advanced selection methods. The results presented here demonstrate the capability of *Euclid* slitless spectroscopy in building large quasar samples and provide a basis for extending this approach to the wide-area AGN censuses enabled by forthcoming *Euclid* data releases.

Data availability

The full catalogue of the identified quasars Table B.1 and the composite spectra (Table 2) are available at the CDS via <https://cdsarc.cds.unistra.fr/viz-bin/cat/J/A+A/711/A285>

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References

- Aihara, H., Arimoto, N., Armstrong, R., et al. 2018, *PASJ*, 70, S4
- Alonso, D., Fabbian, G., Storey-Fisher, K., et al. 2023, *JCAP*, 11, 043
- Assef, R. J., Stern, D., Noirot, G., et al. 2018, *ApJS*, 234, 23
- Bagley, M. B., Finkelstein, S. L., Koekemoer, A. M., et al. 2023, *ApJ*, 946, L12
- Bañados, E., Le Brun, V., Belladitta, S., et al. 2025, *MNRAS*, 542, 1088
- Bertin, E., Schefer, M., Apostolakis, N., et al. 2020, *ASP Conf. Ser.*, 527, 461
- Blanton, M. R., Bershad, M. A., Abolfathi, B., et al. 2017, *AJ*, 154, 28
- Boroson, T. A., & Meyers, K. A. 1992, *ApJ*, 397, 442
- Brammer, G. 2019, Astrophysics Source Code Library [record ascl:1905.001]
- Carnerero, M. I., Raiteri, C. M., Rimoldini, L., et al. 2023, *A&A*, 674, A24
- Chambers, K. C., Magnier, E. A., Metcalfe, N., et al. 2016, ArXiv e-prints [arXiv:1612.05560]
- Chaussidon, E., Yèche, C., Palanque-Delabrouille, N., et al. 2023, *ApJ*, 944, 107
- Cioni, M. R. L., Clementini, G., Girardi, L., et al. 2011, *A&A*, 527, A116
- Conselice, C. J. 2003, *ApJS*, 147, 1
- Conselice, C. J. 2014, *ARA&A*, 52, 291
- Croom, S. M., Smith, R. J., Boyle, B. J., et al. 2004, *MNRAS*, 349, 1397
- De Angeli, F., Weiler, M., Montegriffo, P., et al. 2023, *A&A*, 674, A2
- Delchambre, L., Bailer-Jones, C. A. L., Bellas-Velidis, I., et al. 2023, *A&A*, 674, A31
- DESI Collaboration (Abdul-Karim, M., et al.) 2025, ArXiv e-prints [arXiv:2503.14745]
- Doyon, R., Willott, C. J., Hutchings, J. B., et al. 2023, *PASP*, 135, 098001
- Ducourant, C., Krone-Martins, A., Galluccio, L., et al. 2023, *A&A*, 674, A11
- Edge, A., Sutherland, W., Kuijken, K., et al. 2013, *Messenger*, 154, 32
- Emerson, J., McPherson, A., & Sutherland, W. 2006, *Messenger*, 126, 41
- Estrada, N., Mercurio, A., Vulcani, B., et al. 2023, *A&A*, 671, A146
- Estrada-Carpenter, V., Papovich, C., Momcheva, I., et al. 2019, *ApJ*, 870, 133
- Euclid Collaboration (Lusso, E., et al.) 2024, *A&A*, 685, A108
- Euclid Collaboration (Cropper, M., et al.) 2025, *A&A*, 697, A2
- Euclid Collaboration (Jahnke, K., et al.) 2025, *A&A*, 697, A3
- Euclid Collaboration (Mellier, Y., et al.) 2025, *A&A*, 697, A1
- Euclid Collaboration (Aussel, H., et al.) 2026, *A&A*, 711, A1
- Euclid Collaboration (Copin, Y., et al.) 2026, *A&A*, 711, A6
- Euclid Collaboration (Margalef-Bentabol, B., et al.) 2026, *A&A*, 711, A18
- Euclid Collaboration (Matamoros Zatarain, T., et al.) 2026, *A&A*, 711, A20
- Euclid Collaboration (Quilley, L., et al.) 2026, *A&A*, 711, A8
- Euclid Collaboration (Romelli, E., et al.) 2026, *A&A*, 711, A4
- Euclid Collaboration (Roster, W., et al.) 2026, *A&A*, 711, A16
- Euclid Collaboration (Tarsitano, F., et al.) 2026, *A&A*, 711, A19
- Euclid Collaboration (Walmsley, M., et al.) 2026, *A&A*, 711, A9
- Euclid Quick Release Q1. 2025, <https://doi.org/10.57780/esa-2853f3b>
- Evans, I. N., Evans, J. D., Martínez-Galarza, J. R., et al. 2024, *ApJS*, 274, 22
- Fan, X., Bañados, E., & Simcoe, R. A. 2023, *ARA&A*, 61, 373
- Fitzpatrick, M. J., Olsen, K., Economou, F., et al. 2014, *SPIE Conf. Ser.*, 9149, 91491T
- Flesch, E. W. 2023, *Open J. Astrophys.*, 6, 49
- Fruchter, A. S., & Hook, R. N. 2002, *PASP*, 114, 144
- Fu, Y. 2025, <https://doi.org/10.5281/zenodo.15571037>
- Fu, Y. 2026, <https://doi.org/10.5281/zenodo.18642758>
- Fu, Y., Wu, X.-B., Yang, Q., et al. 2021, *ApJS*, 254, 6
- Fu, Y., Wu, X.-B., Li, Y., et al. 2024, *ApJS*, 271, 54
- Fu, Y., Wu, X.-B., Bouwens, R. J., et al. 2025, *ApJS*, 279, 54
- Gaia Collaboration (Klioner, S. A., et al.) 2022, *A&A*, 667, A148
- Gaia Collaboration (Vallenari, A., et al.) 2023a, *A&A*, 674, A1
- Gaia Collaboration (Bailer-Jones, C. A. L., et al.) 2023b, *A&A*, 674, A41
- Glikman, E., Helfand, D. J., & White, R. L. 2006, *ApJ*, 640, 579
- Gordon, K. 2024, *J. Open Source Softw.*, 9, 7023
- Gordon, K. D., Clayton, G. C., Declair, M., et al. 2023, *ApJ*, 950, 86
- Green, G. M. 2018, *J. Open Source Softw.*, 3, 695
- Hall, P. B., Anderson, S. F., Strauss, M. A., et al. 2002, *ApJS*, 141, 267
- International Organization for Standardization & International Commission on Illumination. 2019, ISO/CIE 11664-4:2019 Colorimetry – Part 4: CIE 1976 L*a*b* Colour Space, <https://www.iso.org/standard/74166.html>
- Jauzac, M., Jullo, E., Kneib, J.-P., et al. 2012, *MNRAS*, 426, 3369
- Juneau, S., Jacques, A., Pothier, S., et al. 2025, *ASP Conf. Ser.*, 541, 77
- Koekemoer, A. M., Faber, S. M., Ferguson, H. C., et al. 2011, *ApJS*, 197, 36
- Kümmel, M., Walsh, J. R., Pirzkal, N., Kuntschner, H., & Pasquali, A. 2009, *PASP*, 121, 59
- Kümmel, M., Álvarez-Ayllón, A., Bertin, E., et al. 2022, ArXiv e-prints [arXiv:2212.02428]
- Landt, H., Bentz, M. C., Ward, M. J., et al. 2008, *ApJS*, 174, 282
- Landt, H., Elvis, M., Ward, M. J., et al. 2011, *MNRAS*, 414, 218
- Landt, H., Ward, M. J., Peterson, B. M., et al. 2013, *MNRAS*, 432, 113
- Lays, C., Ley, C., Klein, O., Bernard, P., & Licata, L. 2013, *J. Exp. Social Psychol.*, 49, 764
- Mainzer, A., Bauer, J., Grav, T., et al. 2011, *ApJ*, 731, 53
- Marocco, F., Eisenhardt, P. R. M., Fowler, J. W., et al. 2021, *ApJS*, 253, 8
- Mateos, S., Alonso-Herrero, A., Carrera, F. J., et al. 2012, *MNRAS*, 426, 3271
- McMahon, R. G., Banerji, M., Gonzalez, E., et al. 2013, *Messenger*, 154, 35
- Merloni, A., Lamer, G., Liu, T., et al. 2024, *A&A*, 682, A34
- Meyer, R. A., Oesch, P. A., Giovannizzo, E., et al. 2024, *MNRAS*, 535, 1067
- Minniti, D., Lucas, P. W., Emerson, J. P., et al. 2010, *New Astron.*, 15, 433
- Miyazaki, S., Komiyama, Y., Kawanomoto, S., et al. 2018, *PASJ*, 70, S1
- Momcheva, I. G., Brammer, G. B., van Dokkum, P. G., et al. 2016, *ApJS*, 225, 27
- Myers, A. D., Outram, P. J., Shanks, T., et al. 2003, *MNRAS*, 342, 467
- Myers, A. D., Palanque-Delabrouille, N., Prakash, A., et al. 2015, *ApJS*, 221, 27
- Navarro, V., del Rio, S., Angel Diego, M., et al. 2024, *Space Data Management. Studies in Big Data*, 141, 1
- Neveux, R., Burtin, E., de Mattia, A., et al. 2020, *MNRAS*, 499, 210
- Nikutta, R., Fitzpatrick, M., Scott, A., & Weaver, B. A. 2020, *Astron. Comput.*, 33, 100411
- Oesch, P. A., Brammer, G., Naidu, R. P., et al. 2023, *MNRAS*, 525, 2864
- Oke, J. B., & Gunn, J. E. 1983, *ApJ*, 266, 713
- Onken, C. A., Wolf, C., Bessell, M. S., et al. 2024, *PASA*, 41, e061
- Osmer, P. S., & Hewett, P. C. 1991, *ApJS*, 75, 273
- Petter, G. C., Hickox, R. C., Alexander, D. M., et al. 2023, *ApJ*, 946, 27
- Pirzkal, N., Xu, C., Malhotra, S., et al. 2004, *ApJS*, 154, 501
- Planck Collaboration XLVIII. 2016, *A&A*, 596, A109
- Pullen, A. R., Alam, S., & Ho, S. 2015, *MNRAS*, 449, 4326
- Ricci, F., Treister, E., Bauer, F. E., et al. 2022, *ApJS*, 261, 8
- Richards, G. T., Fan, X., Newberg, H. J., et al. 2002, *AJ*, 123, 2945
- Rieke, M. J., Kelly, D., & Horner, S. 2005, *SPIE Conf. Ser.*, 5904, 1
- Rieke, M. J., Kelly, D. M., Misselt, K., et al. 2023, *PASP*, 135, 028001
- Rimoldini, L., Holl, B., Gavras, P., et al. 2023, *A&A*, 674, A14
- Roberts-Borsani, G., Morishita, T., Treu, T., et al. 2022, *ApJ*, 938, L13
- Sartoretti, P., Katz, D., Cropper, M., et al. 2018, *A&A*, 616, A6
- Schlafly, E. F., Meisner, A. M., & Green, G. M. 2019, *ApJS*, 240, 30
- Schmidt, M., Schneider, D. P., & Gunn, J. E. 1986, *ApJ*, 306, 411
- Schneider, D. P., Schmidt, M., & Gunn, J. E. 1999, *AJ*, 117, 40
- Sharon, K., Cerny, C., Rigby, J. R., et al. 2022, ArXiv e-prints [arXiv:2207.05709]
- Simcoe, R. A., Burgasser, A. J., Bernstein, R. A., et al. 2008, *SPIE Conf. Ser.*, 7014, 70140U
- Stern, D., Assef, R. J., Benford, D. J., et al. 2012, *ApJ*, 753, 30
- Storey-Fisher, K., Hogg, D. W., Rix, H.-W., et al. 2024, *ApJ*, 964, 69
- Sun, F., Egami, E., Pirzkal, N., et al. 2023, *ApJ*, 953, 53
- Urrutia, T., Becker, R. H., White, R. L., et al. 2009, *ApJ*, 698, 1095
- Vanden Berk, D. E., Richards, G. T., Bauer, A., et al. 2001, *AJ*, 122, 549
- Webb, N. A., Coriat, M., Traulsen, I., et al. 2020, *A&A*, 641, A136
- Williams, C. C., Tacchella, S., Maseda, M. V., et al. 2023, *ApJS*, 268, 64
- Willott, C. J., Doyon, R., Albert, L., et al. 2022, *PASP*, 134, 025002
- Wright, E. L., Eisenhardt, P. R. M., Mainzer, A. K., et al. 2010, *AJ*, 140, 1868
- Wu, X.-B., Hao, G., Jia, Z., Zhang, Y., & Peng, N. 2012, *AJ*, 144, 49

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** Deceased.

Appendix A: Sources with initially discrepant redshifts in the DESI comparison

We illustrate the 16 initially discrepant sources in the DESI consistency check, defined as $|z_{vi} - z_{DESI}|/(1 + z_{DESI}) > 0.15$ among the 454 quasars in common between our Q1 sample and DESI DR1. For each of the 16 sources, we retrieve the DESI DR1 coadded 1D spectrum and the corresponding redshift (z_{DESI}) from the DESI redshift catalogue using the SPARCL; Juneau et al. 2025) and the Astro Data Lab (Fitzpatrick et al. 2014; Nikutta et al. 2020), and compare them to the *Euclid* RGE spectrum used in our visual inspection. The assessment focuses on whether the main emission features are mutually consistent between the DESI and *Euclid* spectra at the proposed redshift solution.

Figure A.1 shows the DESI and *Euclid* spectra together with the template spectrum evaluated at the final z_{vi} . For 11 sources, both DESI and *Euclid* spectra support our visual redshift solution, and z_{DESI} is inconsistent with the spectra. For the remaining five sources whose visual redshifts are revised after reinspection, the panel titles list the updated z_{vi} together with the initial (incorrect) z_{vi} in parentheses; the updated values are adopted in the final catalogue.

As shown in Fig. A.1, most of the 11 sources for which z_{DESI} is inconsistent with the spectra exhibit an unusually red rest-frame UV continuum. In addition, several of these objects (*Euclid* Q1 object_id: 2663629828654232723, 2671669943665629138, 2725249569662970807, 2695896801645034348, and 2718727659662564020) show broad absorption troughs, including absorption associated with Mg II and other metal transitions, consistent with the unusual low-ionisation broad absorption line (LoBAL; see e.g. Boroson & Meyers 1992; Hall et al. 2002; Urrutia et al. 2009) quasars. Because such strongly reddened and LoBAL quasars are relatively rare in the training and template sets used by automated pipelines, and because their red, observed-frame optical continua often have lower S/N than those of typical blue quasars, automated redshift measurements can be less reliable for these objects.

For the remaining five sources, the main causes of the initial misidentifications of z_{vi} are (i) line-identification degeneracies when only a single prominent feature is present, and (ii) artefacts that complicate the interpretation of the spectrum. In particular, in two cases (object_id: 2689353840663761572, and 2680117565652270034) an emission feature consistent with H α at the correct redshift was initially misidentified as Mg II, leading to spuriously high redshift estimates ($z_{vi} > 4$). In another case (object_id: 2664764876670636991), an emission feature consistent with H β at the correct redshift was initially interpreted as H α , leading to an overestimated redshift. Conversely, for object 2675695970669312051, H α was initially misidentified as H β , leading to an underestimated redshift. Finally, for object 2667540229671259497, the *Euclid* spectrum is strongly affected by artefacts (anomalous spike and bump), likely related to contamination from a nearby source, and the initial redshift estimate was driven by the spike seen in the *Euclid* spectrum.

Appendix B: Description of the spectroscopically identified bright quasar catalogue

The format of the catalogue of spectroscopically identified bright quasars from *Euclid* Q1 is shown in Table B.1. The full catalogue will be available at <https://cdsarc.cds.unistra.fr/>.

Appendix C: Composite spectra generation with 1D drizzle

C.1. Details of the 1D drizzle method

We construct *Euclid* quasar composite spectra on a common rest-frame grid with a constant bin size of $\Delta\lambda = 4 \text{ \AA}$. The mapping preserves the integrated flux density in each bin through a linear, flux-conserving rebinning that partitions every input pixel across the overlapping output bins in proportion to the fractional overlap in wavelength. This procedure is mathematically equivalent to a one-dimensional version of the Drizzle algorithm (Fruchter & Hook 2002) and matches the smallest native rest-frame pixel size at the blue end of our spectra, while mildly oversampling the data at longer wavelengths.

Preprocessing and normalisation. Each spectrum is shifted to the rest frame using the adopted redshift, then converted to rest-frame flux density in F_λ via $F_{\lambda,\text{rest}} = (1 + z) F_{\lambda,\text{obs}}$. The spectra are rebinned onto the common rest-frame grid with the flux-conserving drizzle scheme described below. Each rebinned spectrum is then normalised to the running composite within the spectral overlap so that relative shapes are combined consistently; the same normalisation factor is applied to the spectrum's per-pixel variances. We ignore pixels with non-finite or non-positive flux or variance values.

The 1D drizzle rebinning and variance propagation. Let f_p and σ_p^2 be the flux and variance in input pixel p , whose wavelength interval is $[\lambda_{p-1/2}, \lambda_{p+1/2}]$ with width $\Delta\lambda_p = \lambda_{p+1/2} - \lambda_{p-1/2}$. The j -th output bin covers $[\lambda_{j-1/2}, \lambda_{j+1/2}]$ with width $\Delta\lambda_j$. We define the fractional overlap

$$A_{jp} = \frac{\max(0, \min[\lambda_{p+1/2}, \lambda_{j+1/2}] - \max[\lambda_{p-1/2}, \lambda_{j-1/2}])}{\Delta\lambda_p}, \quad (\text{C.1})$$

so that $0 \leq A_{jp} \leq 1$ and $\sum_j A_{jp} = 1$ for fully covered pixels. The rebinned flux and variance for a single spectrum in bin j are then

$$\hat{f}_j = \sum_p A_{jp} f_p, \quad (\text{C.2})$$

$$\hat{\sigma}_j^2 = \sum_p A_{jp}^2 \sigma_p^2, \quad (\text{C.3})$$

assuming uncorrelated input pixels. This operation introduces correlations between neighbouring output bins, but it preserves total flux and provides a well-defined per-bin variance. In the following, we treat $\hat{\sigma}_j$ as the effective per-bin uncertainty and note the presence of correlations where relevant.

Equal-weight mean and its uncertainty. For the mean composite we combine, in each bin j , all rebinned spectra that have valid data in that bin with equal weights. If n_j spectra contribute and their rebinned fluxes and variances are $\hat{f}_{ij}$ and $\hat{\sigma}_{ij}^2$, the mean composite and its formal uncertainty are

$$\bar{f}_j = \frac{1}{n_j} \sum_{i=1}^{n_j} \hat{f}_{ij}, \quad (\text{C.4})$$

$$\sigma_{\bar{f}_j}^2 = \frac{1}{n_j^2} \sum_{i=1}^{n_j} \hat{\sigma}_{ij}^2. \quad (\text{C.5})$$

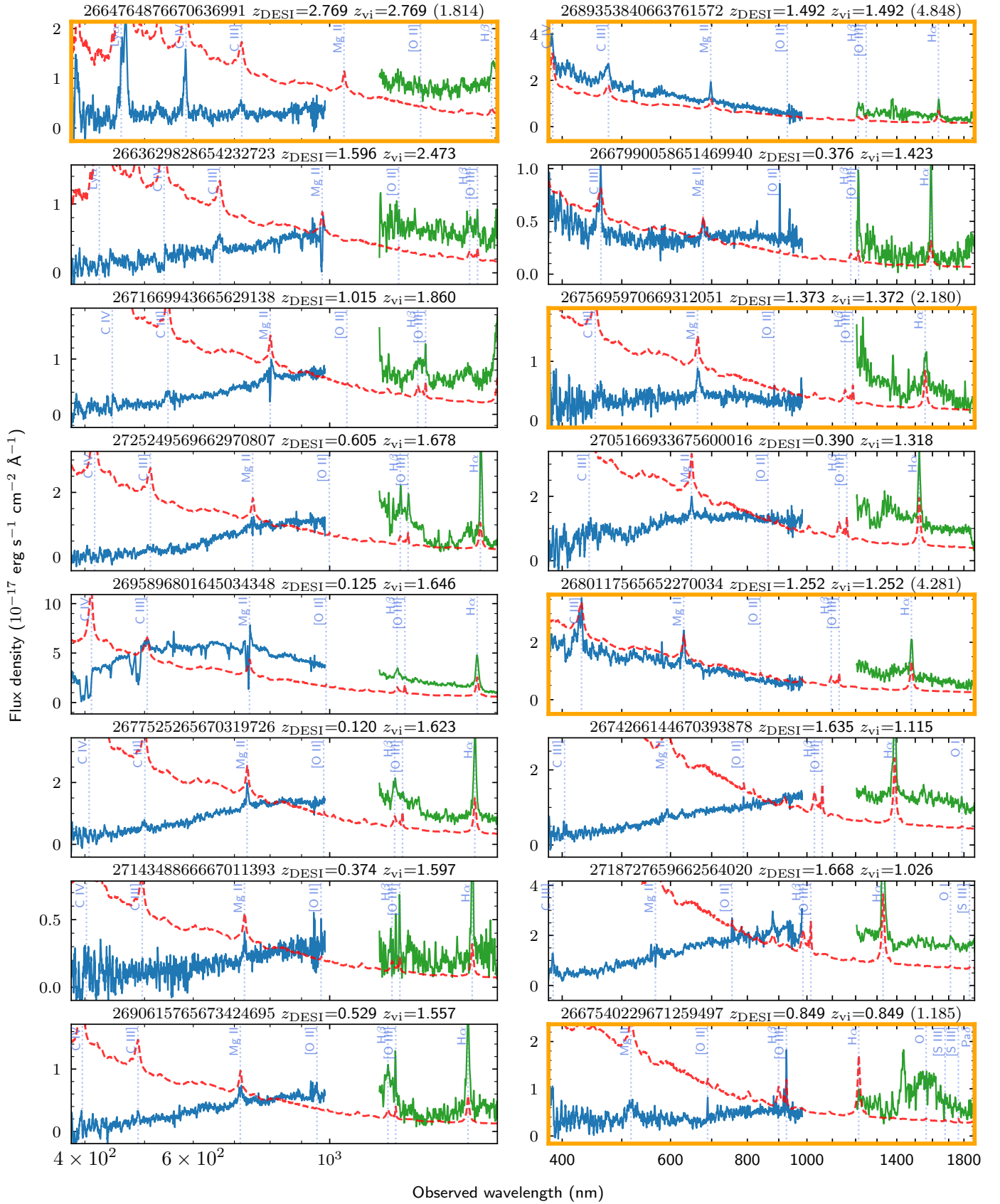


Fig. A.1. DESI (blue) and *Euclid* (green) spectra for the 16 objects that initially satisfy $|z_{\text{vi}} - z_{\text{DESI}}| / (1 + z_{\text{DESI}}) > 0.15$ among the 454 quasars in common with DESI DR1. The dashed red curve shows the quasar template spectrum evaluated at the final z_{vi} (scaled for display). Vertical dotted lines mark the expected observed wavelengths of common quasar emission lines at the final z_{vi} . Panels outlined in orange denote the five sources for which we revise z_{vi} after reinspection; for these, the initial (incorrect) z_{vi} is reported in parentheses after the updated z_{vi} in the panel title.

Table B.1. Format of the catalogue of spectroscopically identified bright quasars from *Euclid* Q1.

Column	Name	Unit	Type	Description
1	object_id	...	long	Euclid Q1 unique source identifier
2	name	...	string	IAU-formatted source name (EUCL JHHMMSS.ss+DDMMSS.s)
3	ra	deg	double	Source barycentre right ascension coordinate in decimal degrees
4	dec	deg	double	Source barycentre declination coordinate in decimal degrees
5	class_vi	...	string	Visual classification of the source
6	z_vi	...	double	Visual redshift of the source
7	sum_mask	...	int	Sum of the mask values of the spectrum within [12047, 18734] Å
8	n_invalid	...	int	Number of invalid pixels in the spectrum within [12047, 18734] Å
9	med_snr	...	double	Median signal-to-noise ratio of the spectrum within [12047, 18734] Å
10	mumax_minus_mag	mag arcsec ⁻²	double	The difference between mu_max and mag_stargal_sep, valid even for NIR-only sources
11	kron_radius	pix	double	Major semi-axis (in pixels) of the elliptical aperture used for total (Kron) aperture photometry on the detection image
12	gaia_id	...	long	The associated <i>Gaia</i> DR3 source id
13	f_psf	...	double	PSF fraction of the source from VIS image (Euclid Collaboration: Margalef-Bentabol et al. 2026)
14	mag_vis_psf	mag	double	VIS I_E band PSF-fitting AB magnitude (from flux_vis_psf)
15	magerr_vis_psf	mag	double	Error on VIS I_E band PSF-fitting AB magnitude
16	mag_y_tmplfit	mag	double	Y_E band template-fit AB magnitude (from flux_y_tmplfit)
17	magerr_y_tmplfit	mag	double	Error on Y_E band template-fit AB magnitude
18	mag_j_tmplfit	mag	double	J_E band template-fit AB magnitude (from flux_j_tmplfit)
19	magerr_j_tmplfit	mag	double	Error on J_E band template-fit AB magnitude
20	mag_h_tmplfit	mag	double	H_E band template-fit AB magnitude (from flux_h_tmplfit)
21	magerr_h_tmplfit	mag	double	Error on H_E band template-fit AB magnitude

Notes. The catalogue described by Table B.1 is available in its entirety at the CDS.

In the implementation we apply sigma clipping in each wavelength bin to reduce the impact of outliers before computing the mean. The sigma-clipping is performed using `astropy.stats.sigma_clip`, with a threshold of $\sigma=6$, and a robust standard deviation estimator calculated as

$$\sigma_R \approx 1.4826 \text{ MAD}, \quad (\text{C.6})$$

where MAD is the median absolute deviation (Leys et al. 2013). For a univariate dataset $X_1, X_2, \dots, X_n$, the MAD is defined as the median of the absolute deviations from the data's median

$$\text{MAD} = \text{median}(|X_i - \bar{X}|). \quad (\text{C.7})$$

In addition, we track the observed object-to-object dispersion in each bin after sigma-clipping via the central second moment

$$s_j^2 = \frac{1}{n_j} \sum_{i=1}^{n_j} (\hat{f}_{ij} - \bar{f}_j)^2, \quad (\text{C.8})$$

whose square root is stored as the ‘RMS’ spectrum. This s_j includes both intrinsic diversity and measurement noise and is quoted separately from the statistical uncertainty $\sigma_{\bar{f}_j}$ on the mean.

Median composite. We formed a median composite by taking, in each bin, the median of the contributing $\hat{f}_{ij}$ values. We applied the same sigma clipping as done when computing the mean composite. We did not attach formal uncertainties to the median composite.

Geometric-mean composite for continuum shape. To preserve the average continuum slope of quasars, we also computed a geometric-mean composite on the same rebinned spectra and wavelength grid, again using equal weights. For contributors with positive flux,

$$\mu_j = \frac{1}{n_j} \sum_{i=1}^{n_j} \ln \hat{f}_{ij}, \quad (\text{C.9})$$

$$f_j^{\text{geo}} = \exp(\mu_j). \quad (\text{C.10})$$

Using error propagation for $\ln \hat{f}_{ij}$ with $\text{Var}[\ln \hat{f}_{ij}] \approx \widehat{\sigma}_{ij}^2 / \hat{f}_{ij}^2$ when $\widehat{\sigma}_{ij} \ll \hat{f}_{ij}$, the uncertainty on μ_j and the geometric-mean flux are

$$\sigma_{\mu_j}^2 = \frac{1}{n_j^2} \sum_{i=1}^{n_j} \frac{\widehat{\sigma}_{ij}^2}{\hat{f}_{ij}^2}, \quad (\text{C.11})$$

$$\sigma_j^{\text{geo}} = f_j^{\text{geo}} \sigma_{\mu_j}. \quad (\text{C.12})$$

We masked bins where fewer than three spectra contribute or where negative flux values prevent a robust logarithmic mean from being determined.

Effect of wavelength bin size. In Sect. 4.3 we adopted a bin size of $\Delta\lambda = 4 \text{ Å}$ for constructing the *Euclid* mean composite spectrum. To assess the impact of this choice, we recomputed the mean composite using coarser grids with $\Delta\lambda = 8 \text{ Å}$ and 13.4 Å , keeping all other steps of the stacking procedure fixed. As shown in Fig. C.1, the continuum shapes of three mean composites are virtually identical in both the optical and near-infrared ranges, while the two larger bin sizes produce undersampled, flattened emission line peaks. This test demonstrates that the bin size of

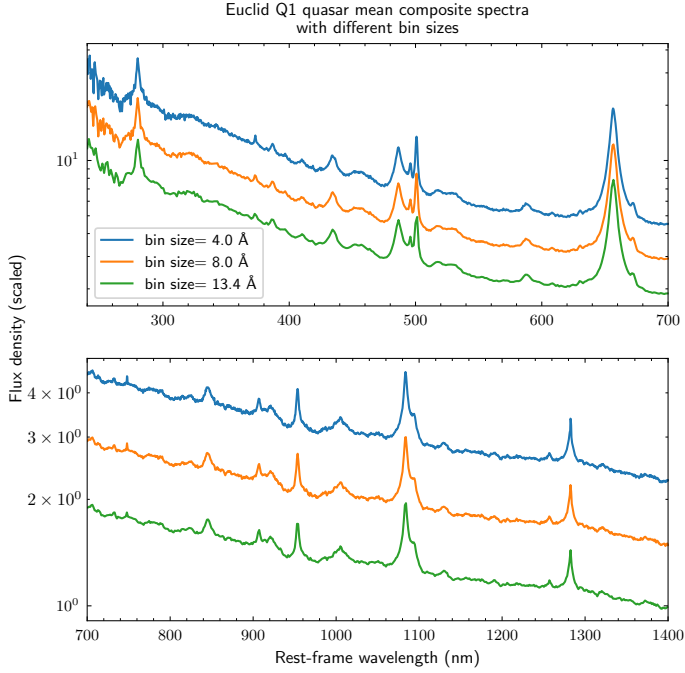


Fig. C.1. Mean *Euclid* Q1 quasar composite spectra constructed with different rest-frame wavelength bin sizes. The three curves show mean composites built on grids with $\Delta\lambda = 4.0 \text{ Å}$ (blue), 8.0 Å (orange), and 13.4 Å (green). For clarity, the spectra are vertically offset by multiplicative factors, as indicated on the right-hand side. The top panel displays the optical range (300–700 nm) and the bottom panel shows the near-infrared range (700–1400 nm). The continua of the three spectra show close agreement, while the two spectra with larger bin sizes have flattened emission line peaks.

$\Delta\lambda = 4 \text{ Å}$ is favourable for the preservation of details of the composite spectra.

C.2. Construction of the piecewise quasar template

The final rest-frame template combines

1. [Vanden Berk et al. \(2001\)](#) for $\lambda_{\text{rest}} < 320 \text{ nm}$;
2. The mean composite from this work for $320 \leq \lambda_{\text{rest}} \leq 1550 \text{ nm}$; and
3. [Glikman et al. \(2006\)](#) for $\lambda_{\text{rest}} > 1550 \text{ nm}$.

All three spectra are resampled onto a common rest-frame grid from 90 to 2100 nm with a spacing of 0.2 nm using a flux-conserving resampler (`specutils.FluxConservingResampler`). Before splicing, we scale the adjoining segments using the mean flux in 5 nm windows centred on each join. At 320 nm we scale our composite to match [Vanden Berk et al. \(2001\)](#); at 1550 nm we scale [Glikman et al. \(2006\)](#) to match our composite. If $\langle F_A \rangle$ and $\langle F_B \rangle$ are the mean fluxes in the overlap window, the scale factor applied to F_B is $s = \langle F_A \rangle / \langle F_B \rangle$. After scaling, we combine the spectra piecewise at the join wavelengths (no additional cross-fade). The resulting template covers 90 to 2100 nm on a uniform grid.

Appendix D: Example cutouts of quasars

We present the *Euclid* spectra and image cutouts of a sample of low-redshift ($z < 0.5$) quasars in Fig. D.1, and a sample of intermediate-redshift ($0.5 < z < 2$) quasars in Fig. D.2.

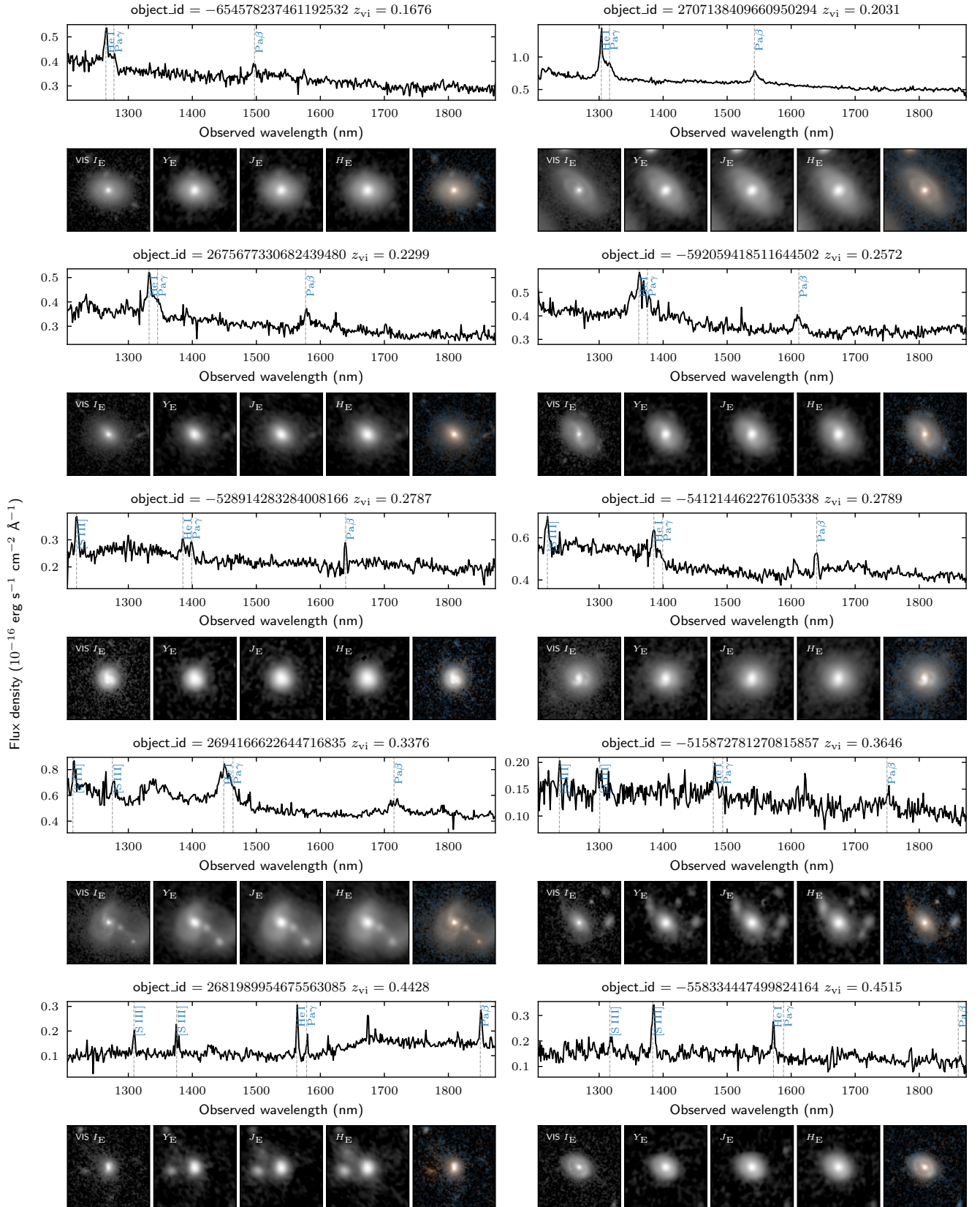


Fig. D.1. *Euclid*1D spectra and imaging cutouts for a random sample of 10 low-redshift ($z < 0.5$) sources. Each panel displays the 1D spectrum (top) and five imaging cutouts with $10''$ sizes (bottom) in the I_E , Y_E , J_E , and H_E bands, as well as a VIS- Y_E composite. The composite image is generated by mapping the VIS and Y_E fluxes into the blue and red channels, respectively, with their mean used for green, and the VIS band used to define overall luminosity in the $L^*a^*b^*$ colour space (International Organization for Standardization & International Commission on Illumination 2019) to enhance morphological detail. Major emission lines detected in the wavelength range [12 047, 18 734] Å are marked.

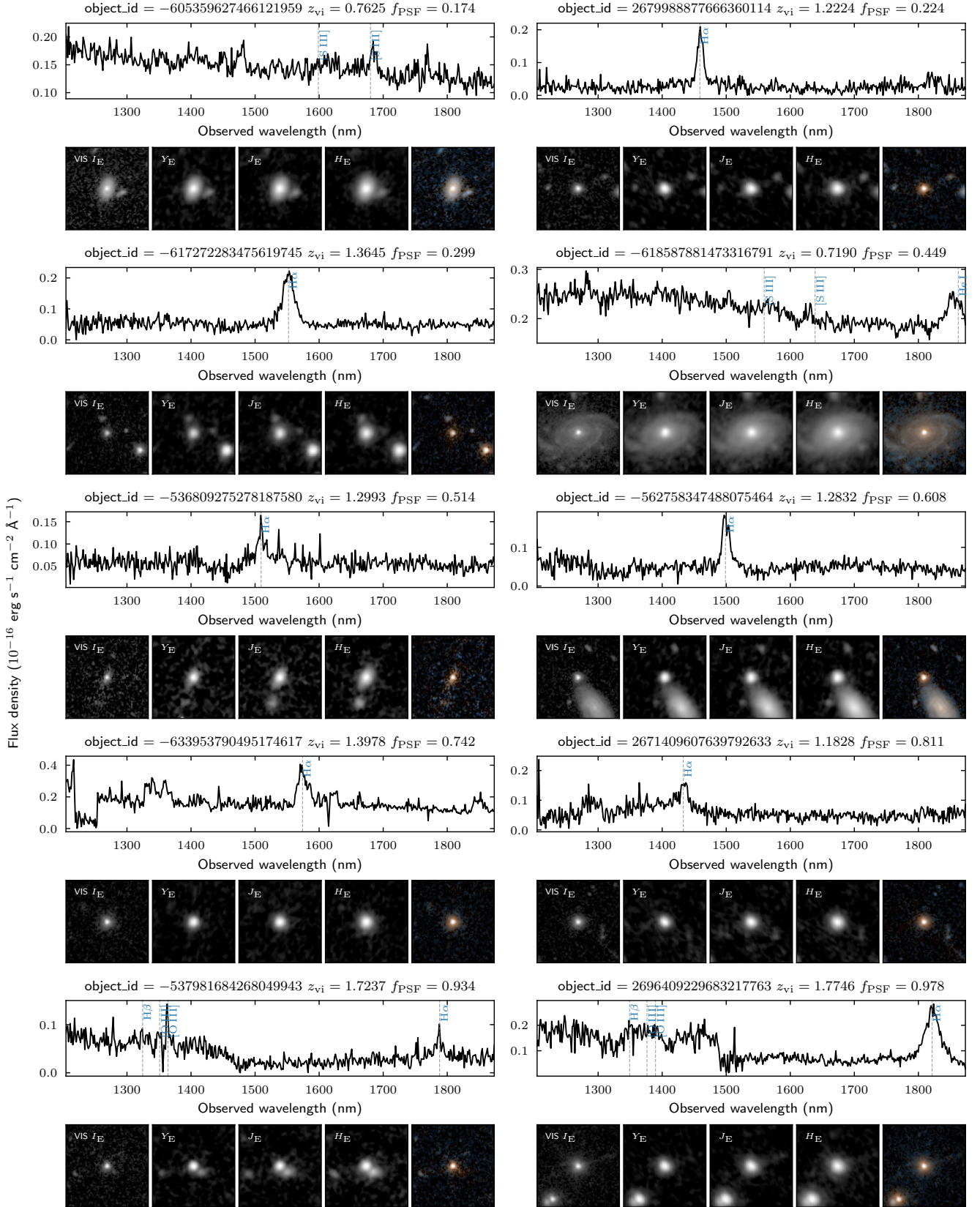


Fig. D.2. Same as Fig. D.1, but for 10 intermediate-redshift ($0.5 < z < 2$) quasars with different f_{PSF} levels.